

Neutrinos (part II): The Ghosts of the Universe

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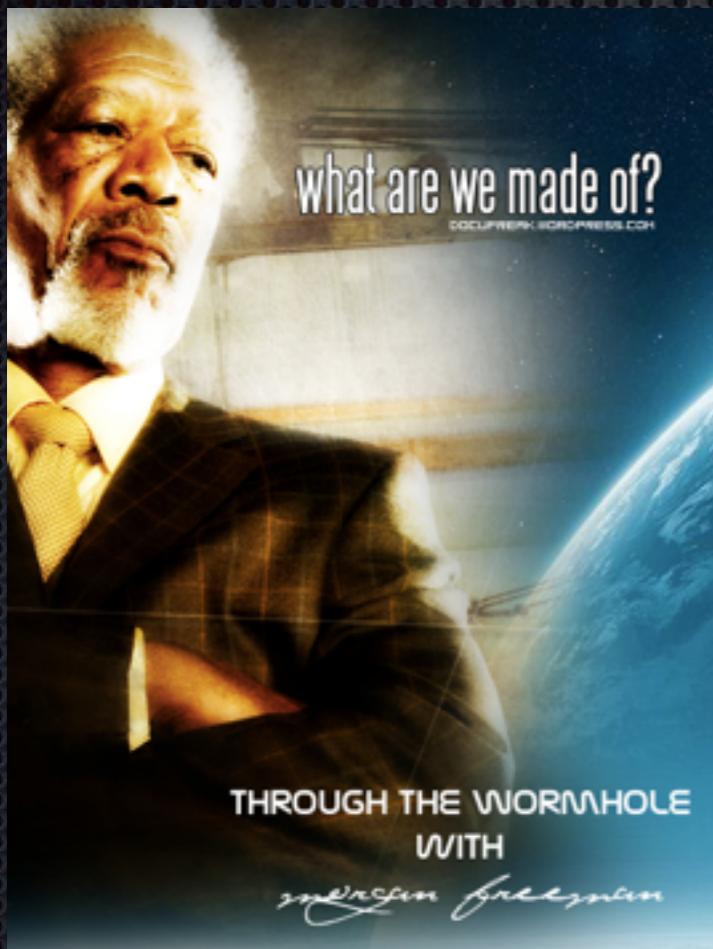
30 June 2015
Fermilab Undergraduate Lecture Series

- ❖ What's particle physics? (translated into “Tia Speak”)
- ❖ Where are we in neutrino physics? (Story time!)
- ❖ FUN MATH for neutrino masses!
- ❖ Where are we going in neutrino physics? (A.K.A why Tia researches neutrinos)

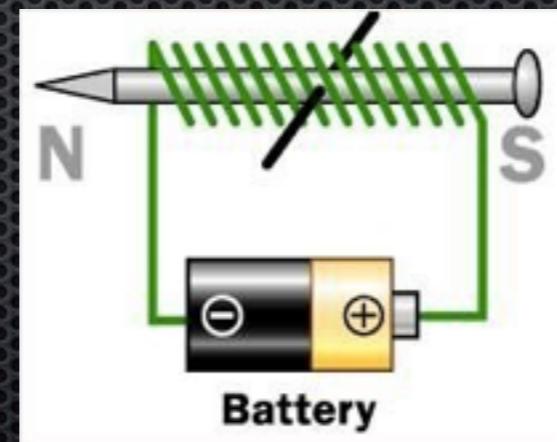
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Particle physics questions

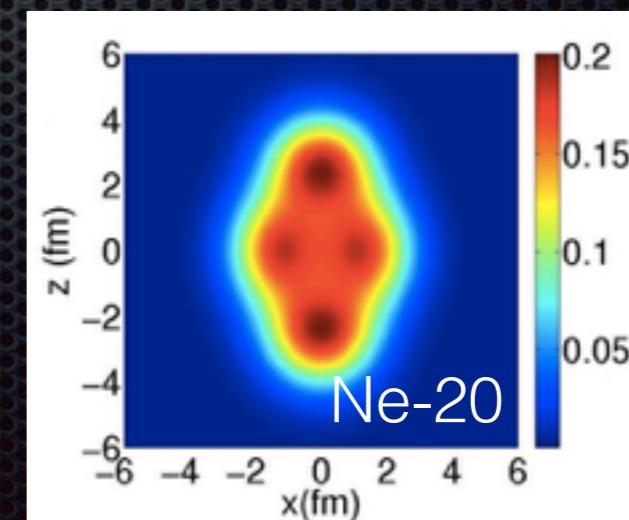
What are we made of?



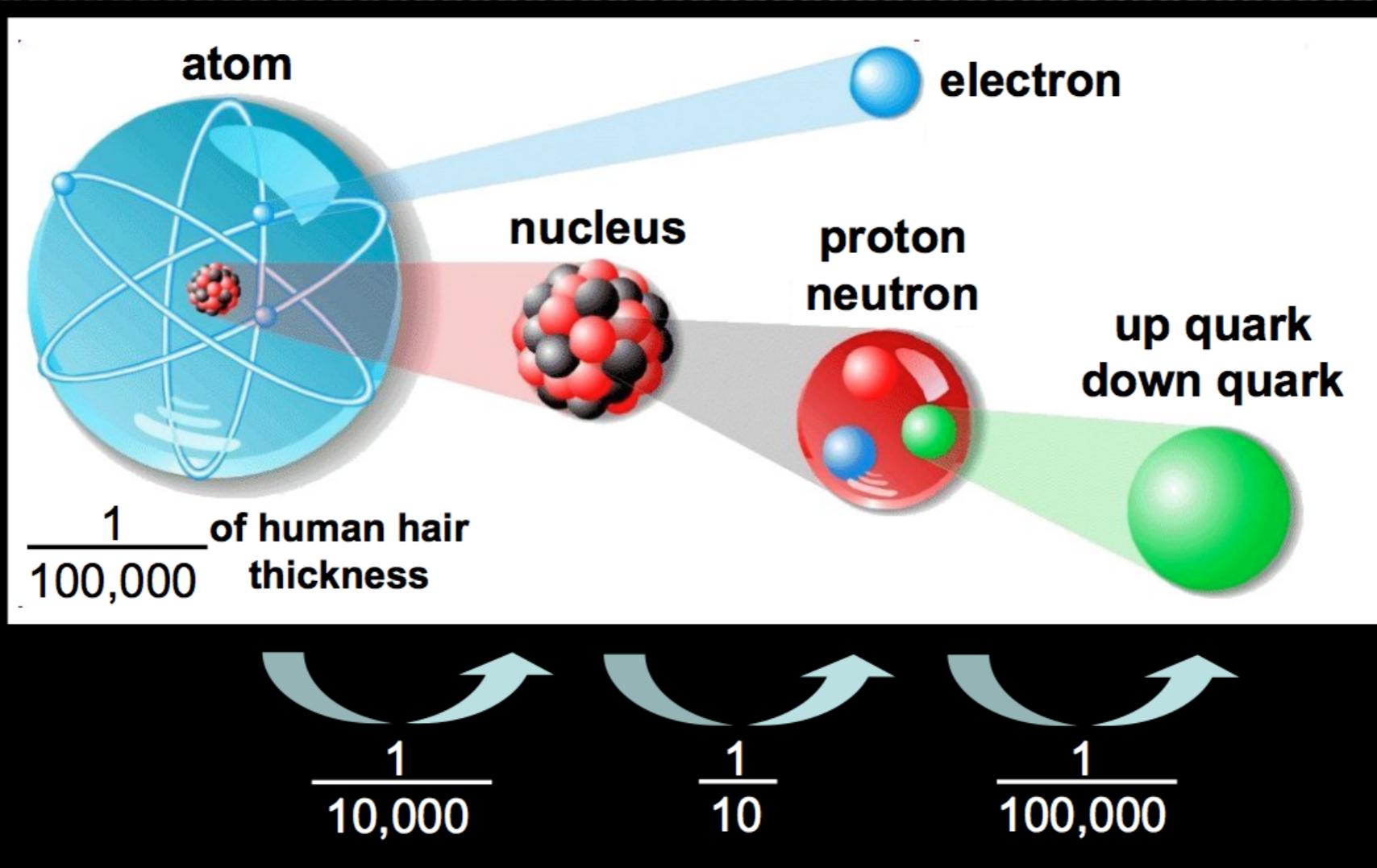
How does it work?



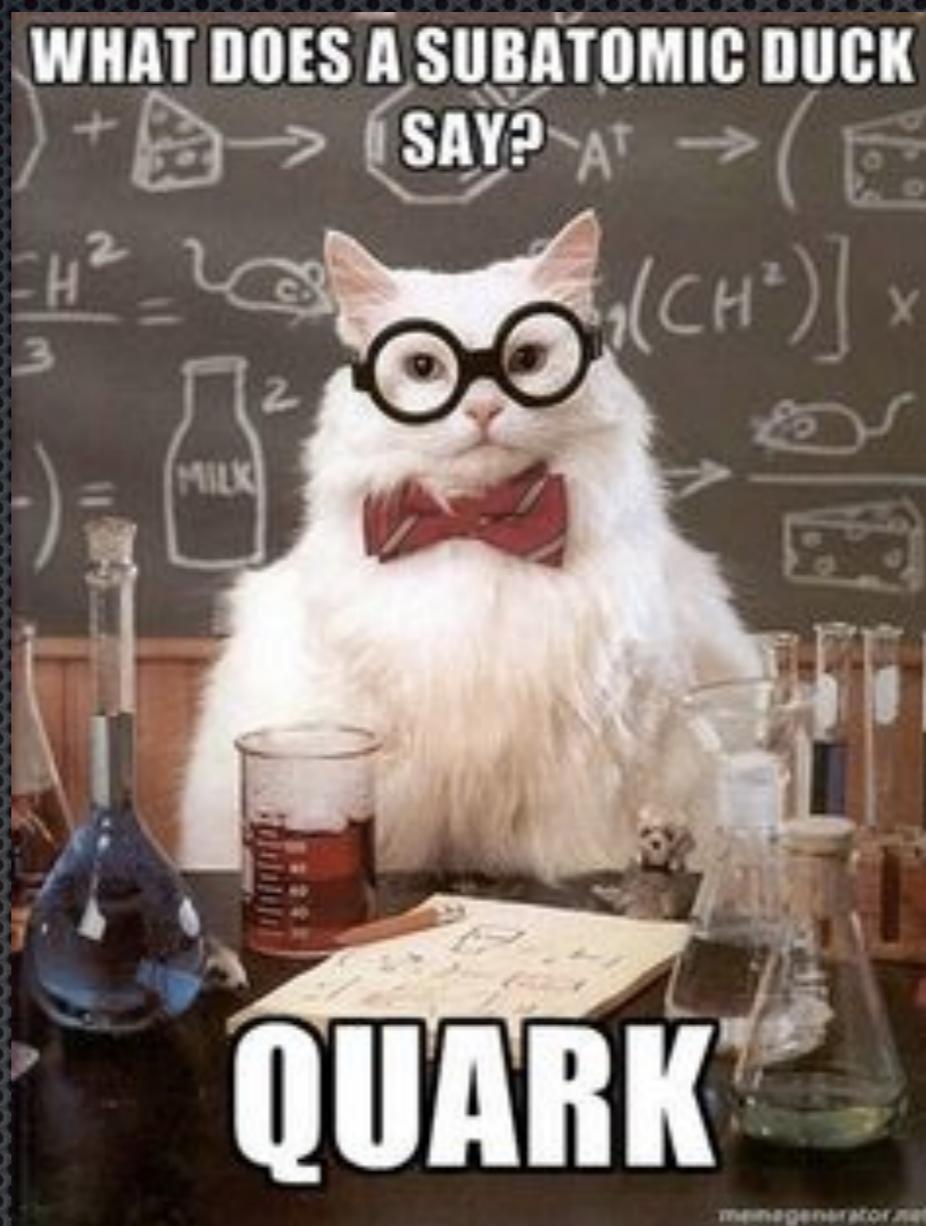
How sure are we?



Particle physics is everywhere

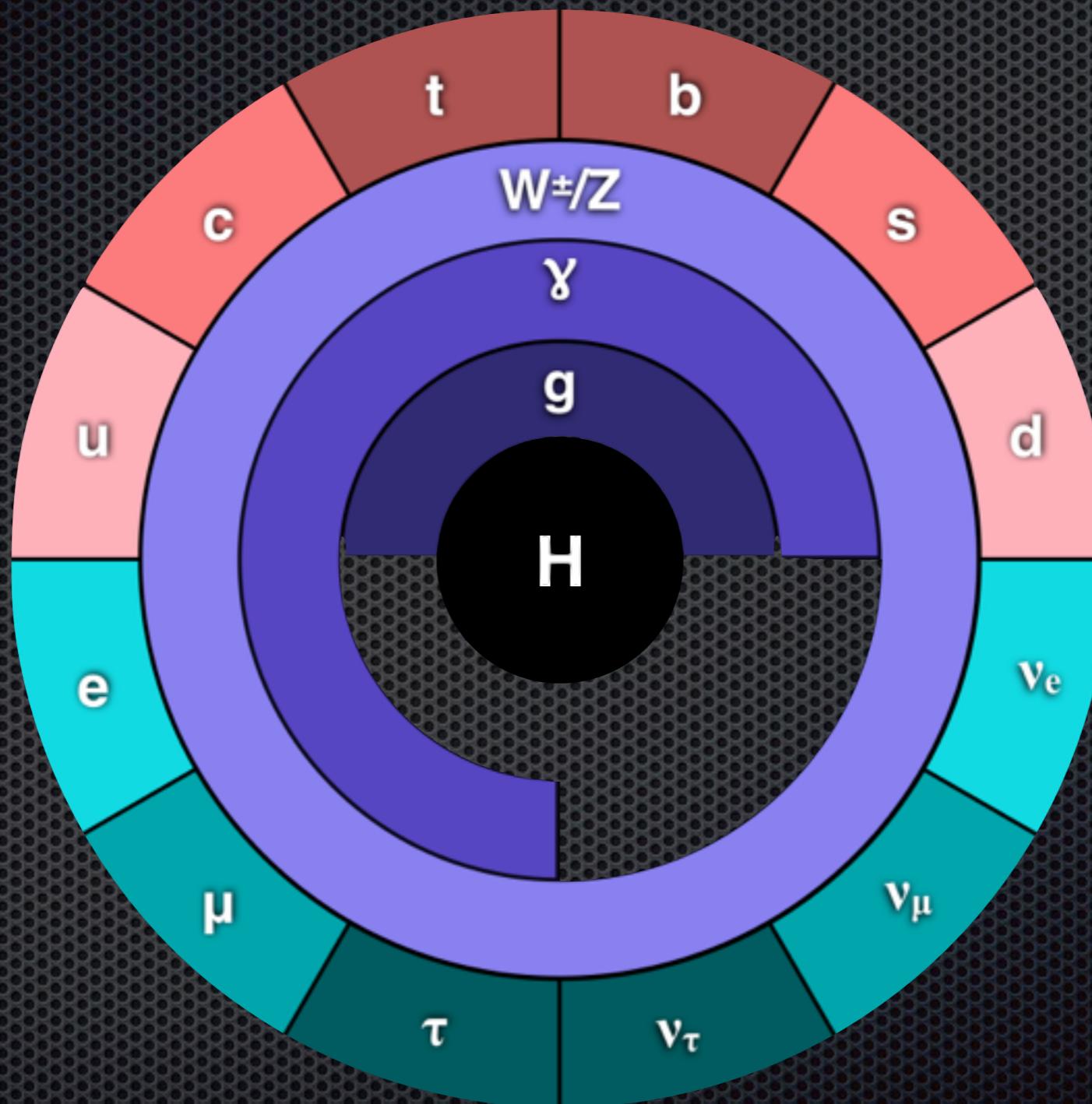


Particle physics is everywhere

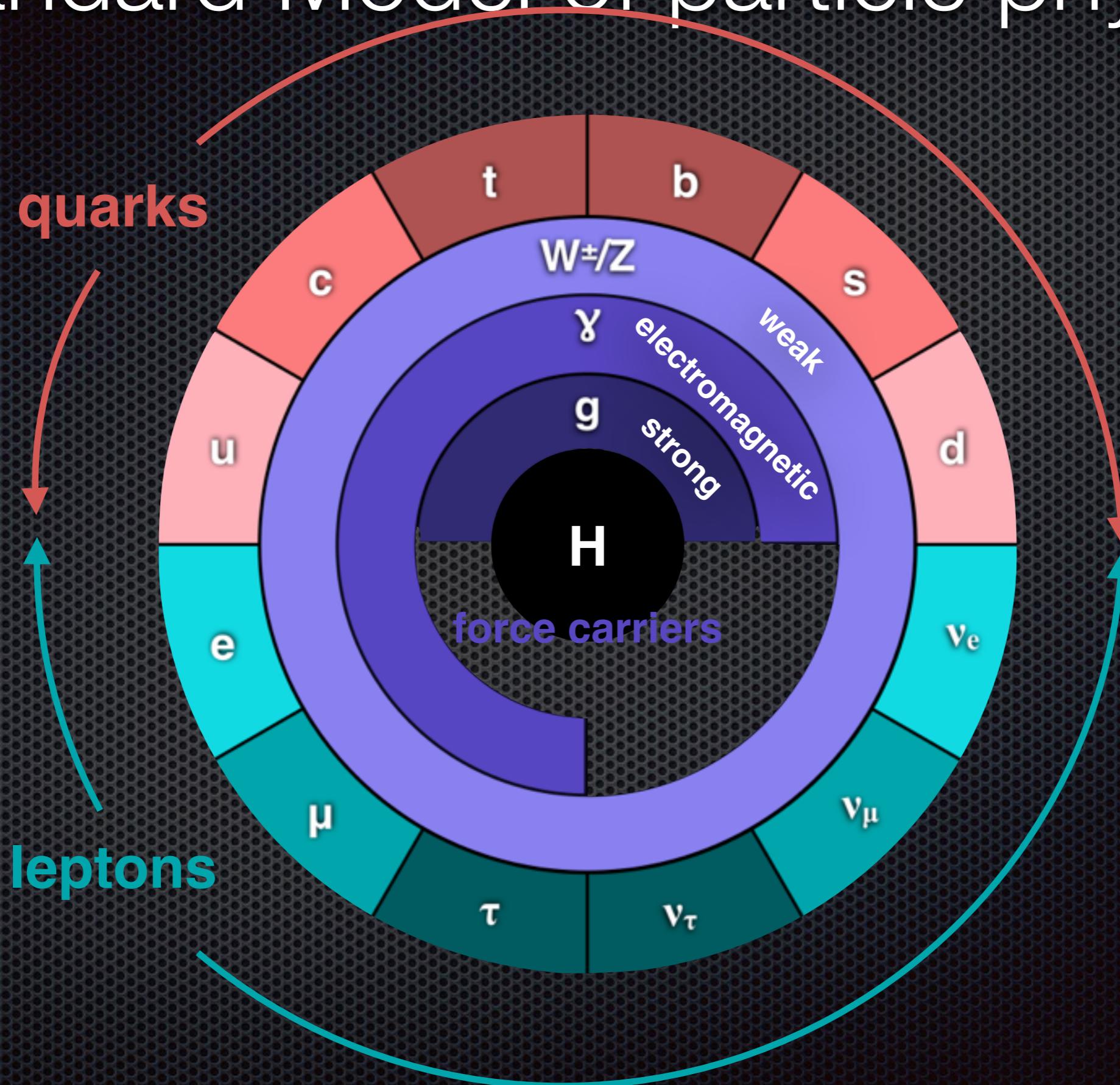


Up quarks, down quarks, and electrons... oh my!
What else?

Standard Model of particle physics

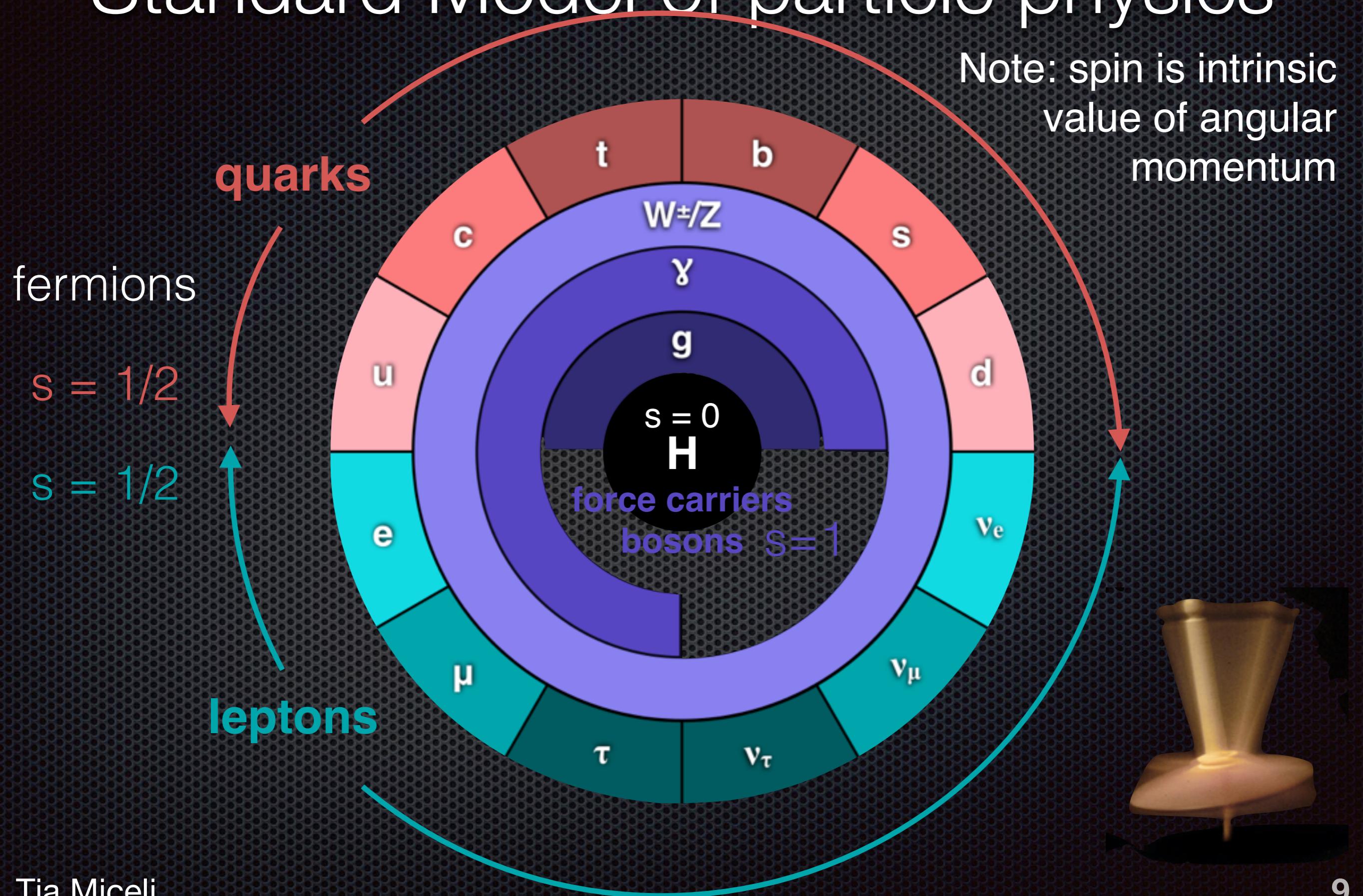


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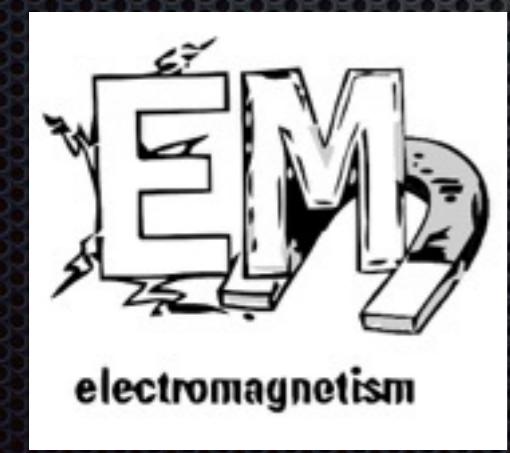
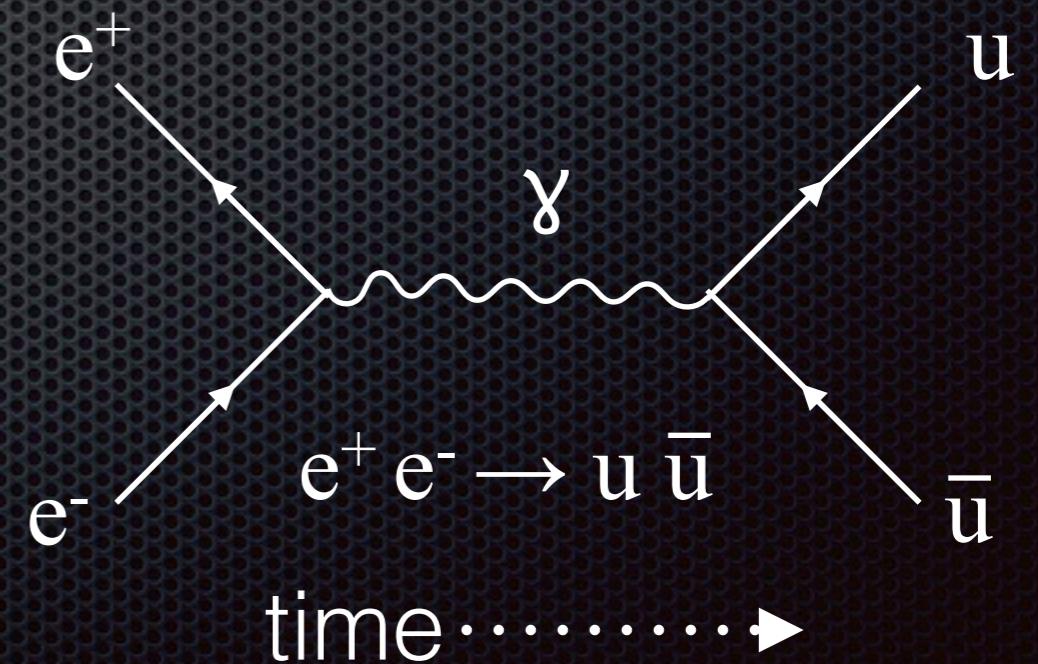
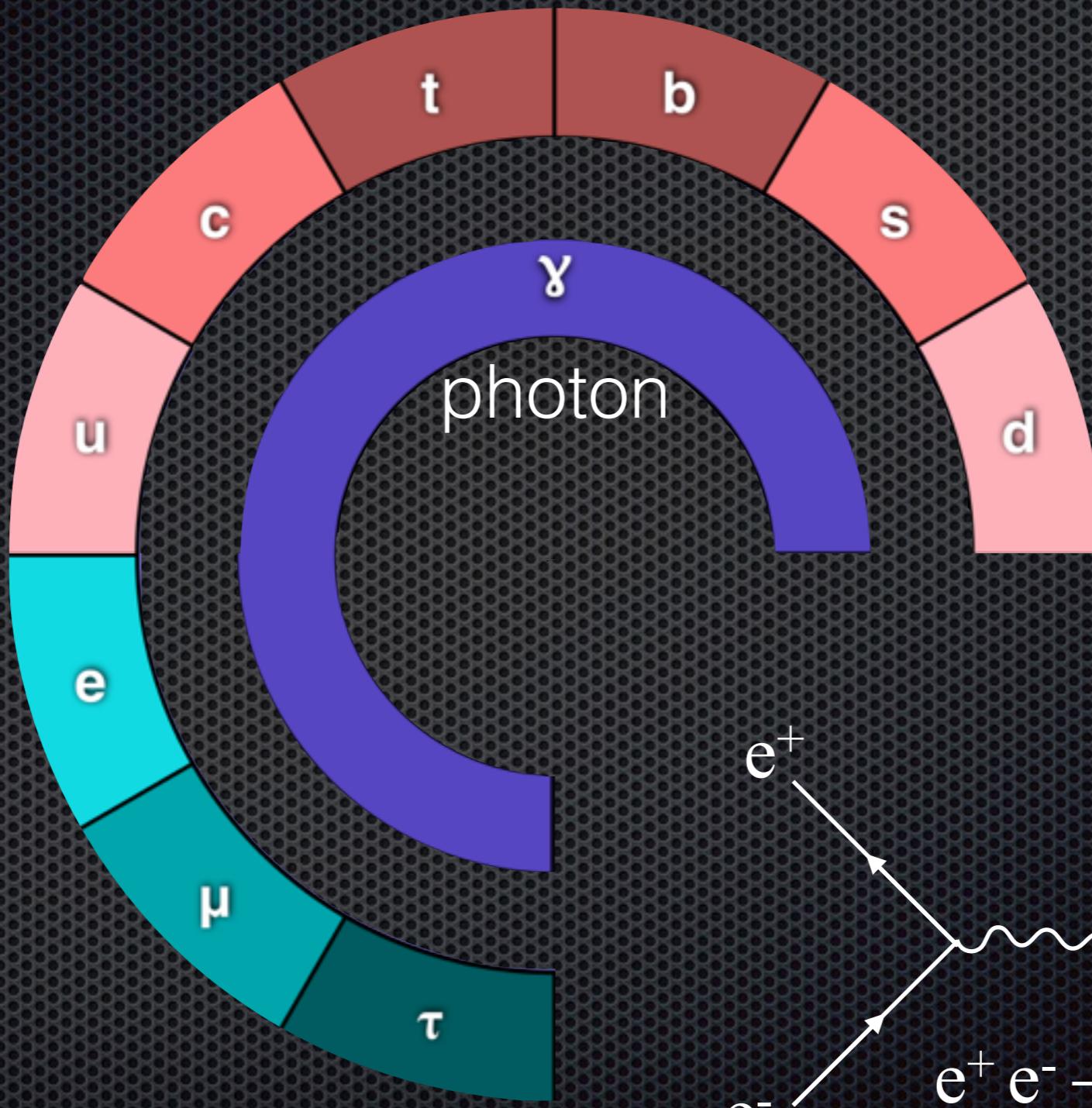


Standard Model of particle physics

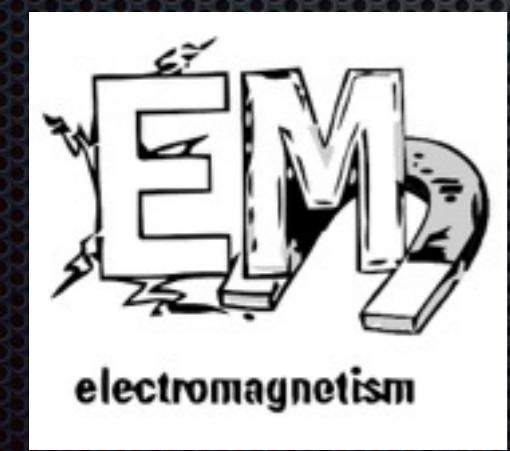
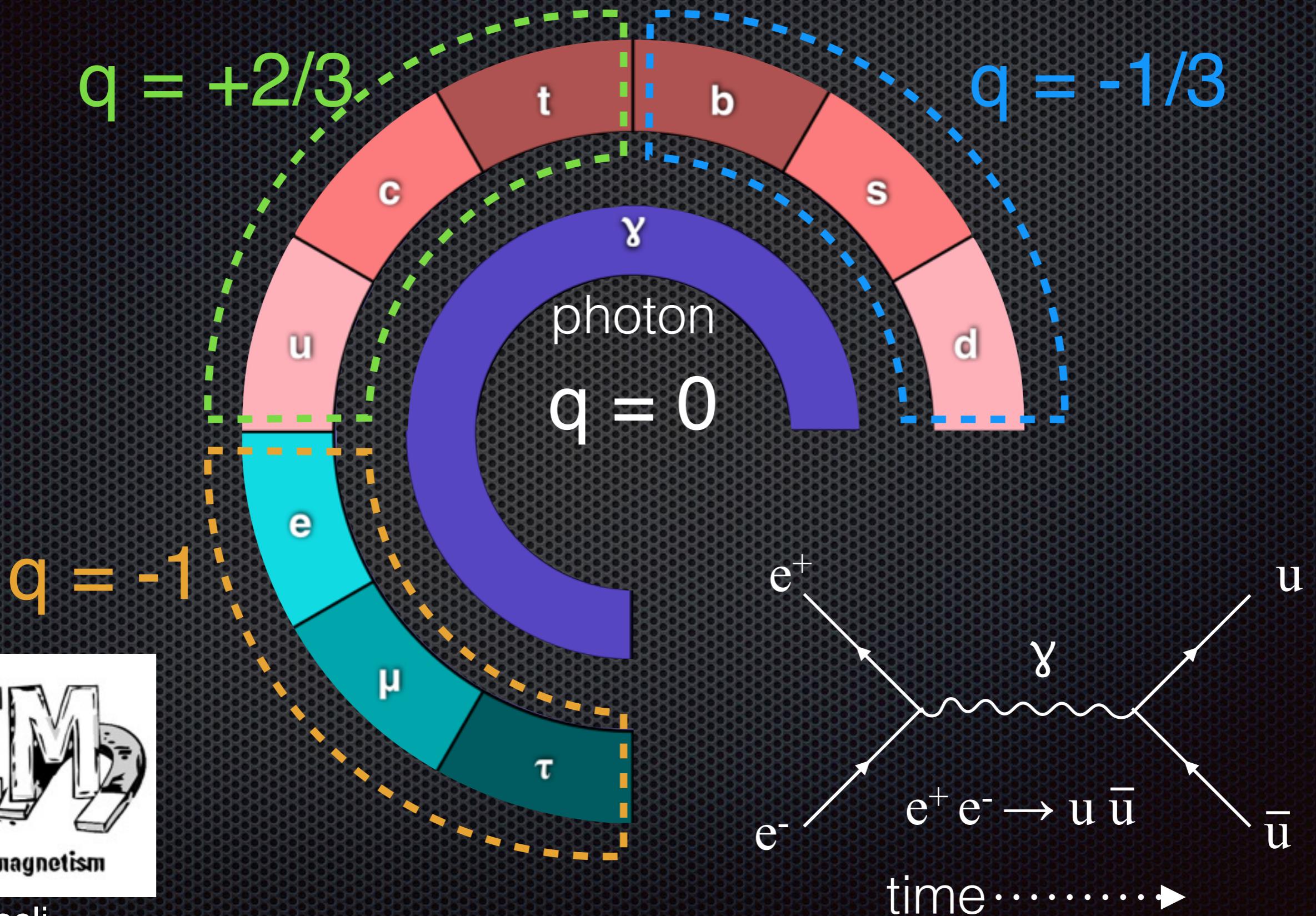
Note: spin is intrinsic value of angular momentum



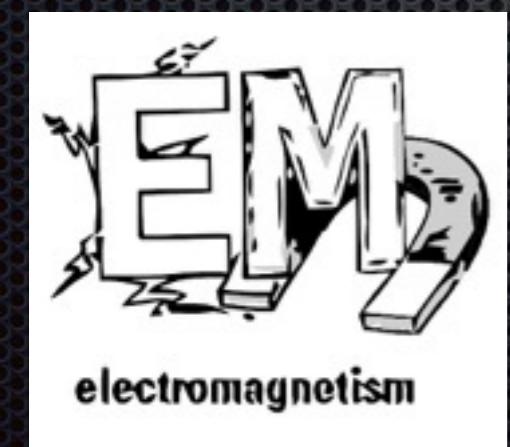
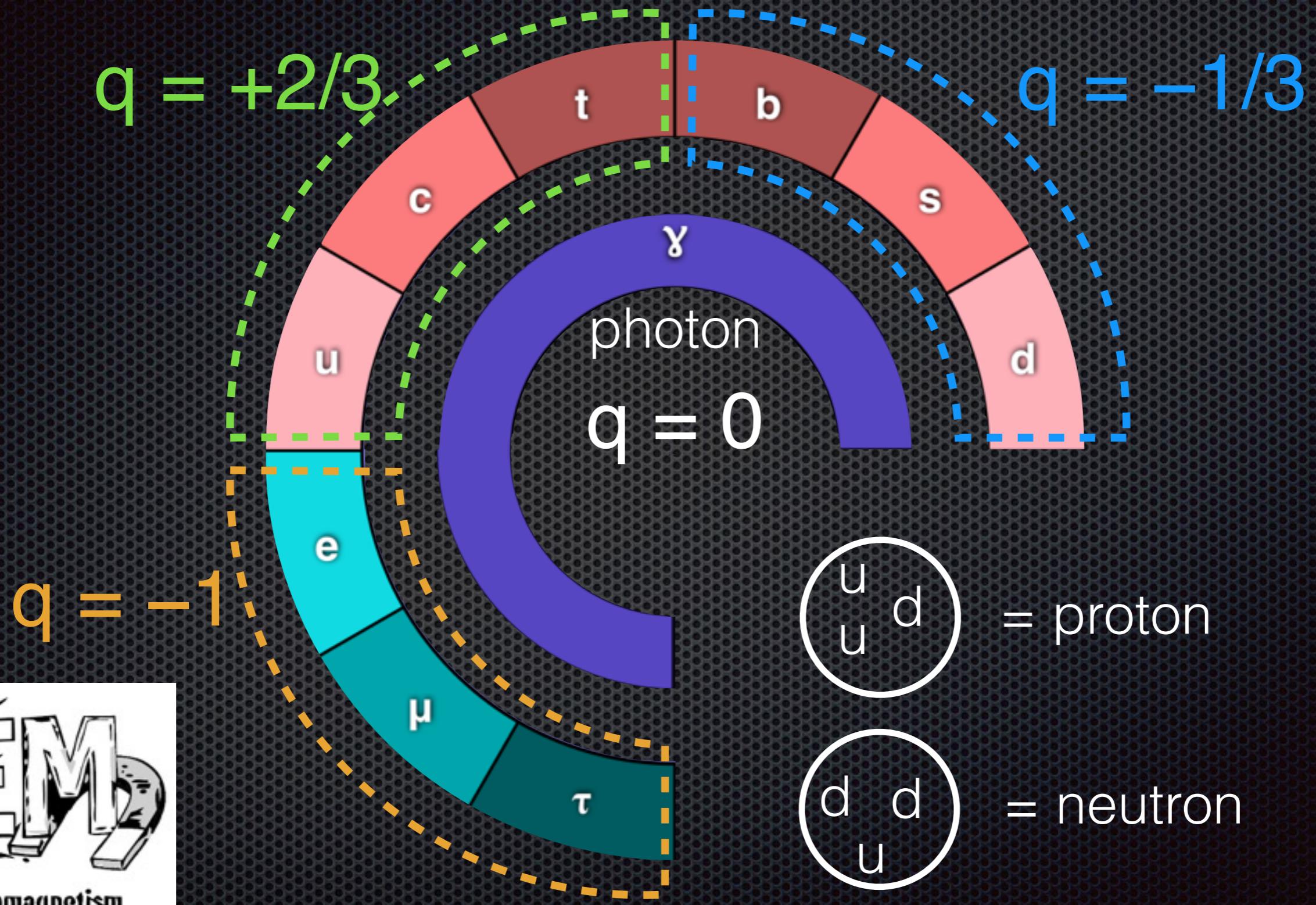
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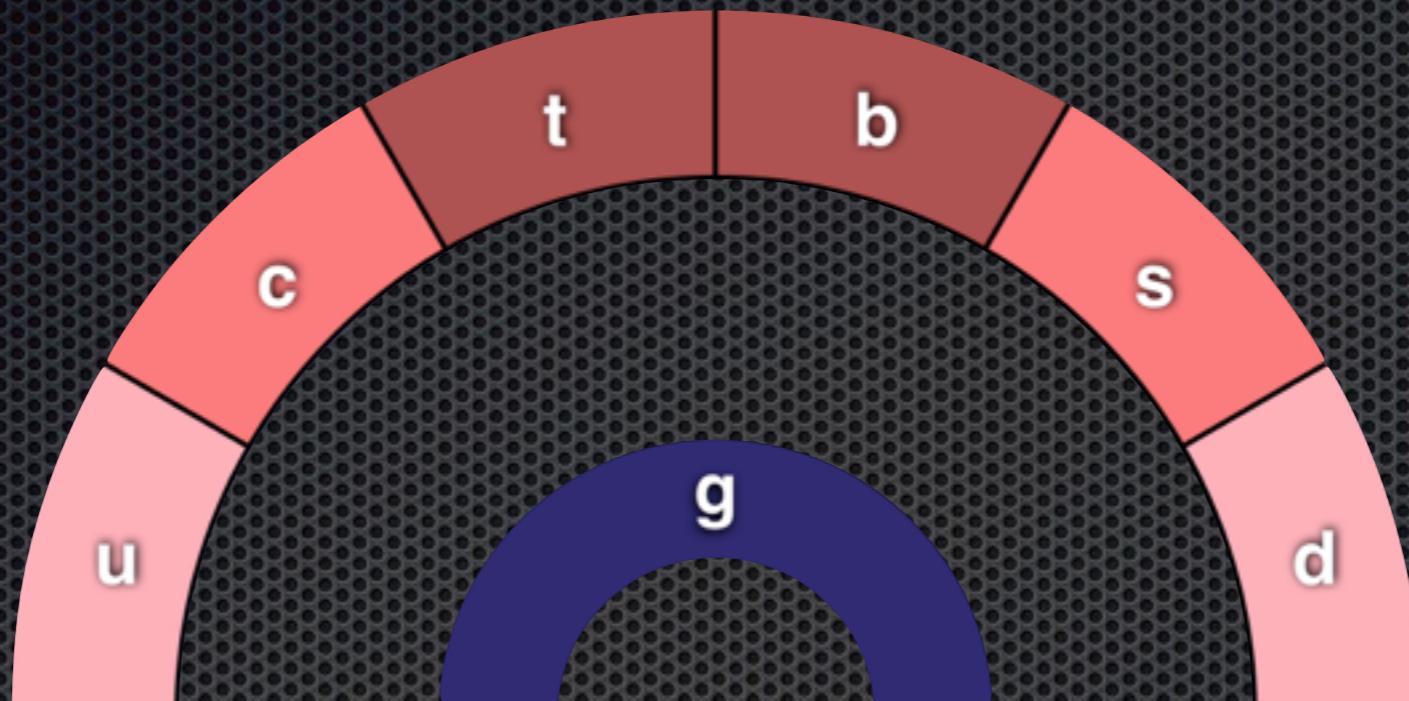
Standard Model of particle physics



Standard Model of particle physics



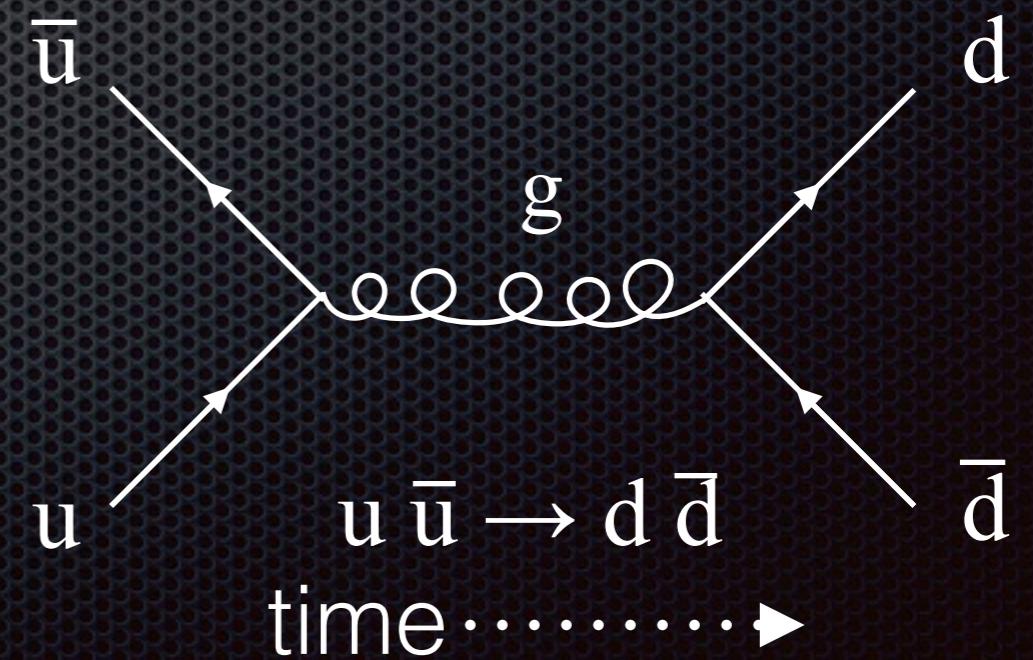
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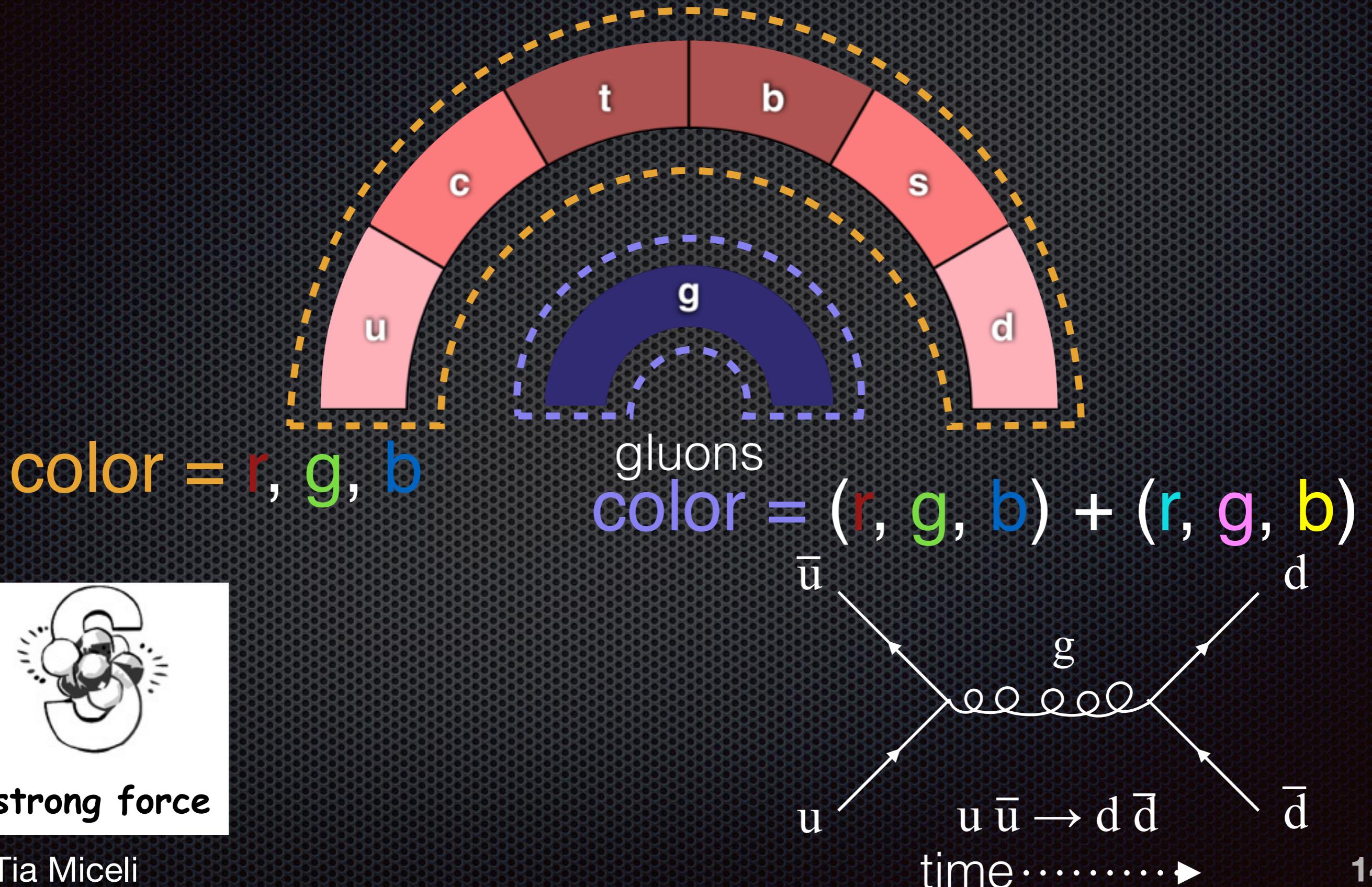
gluons



strong force

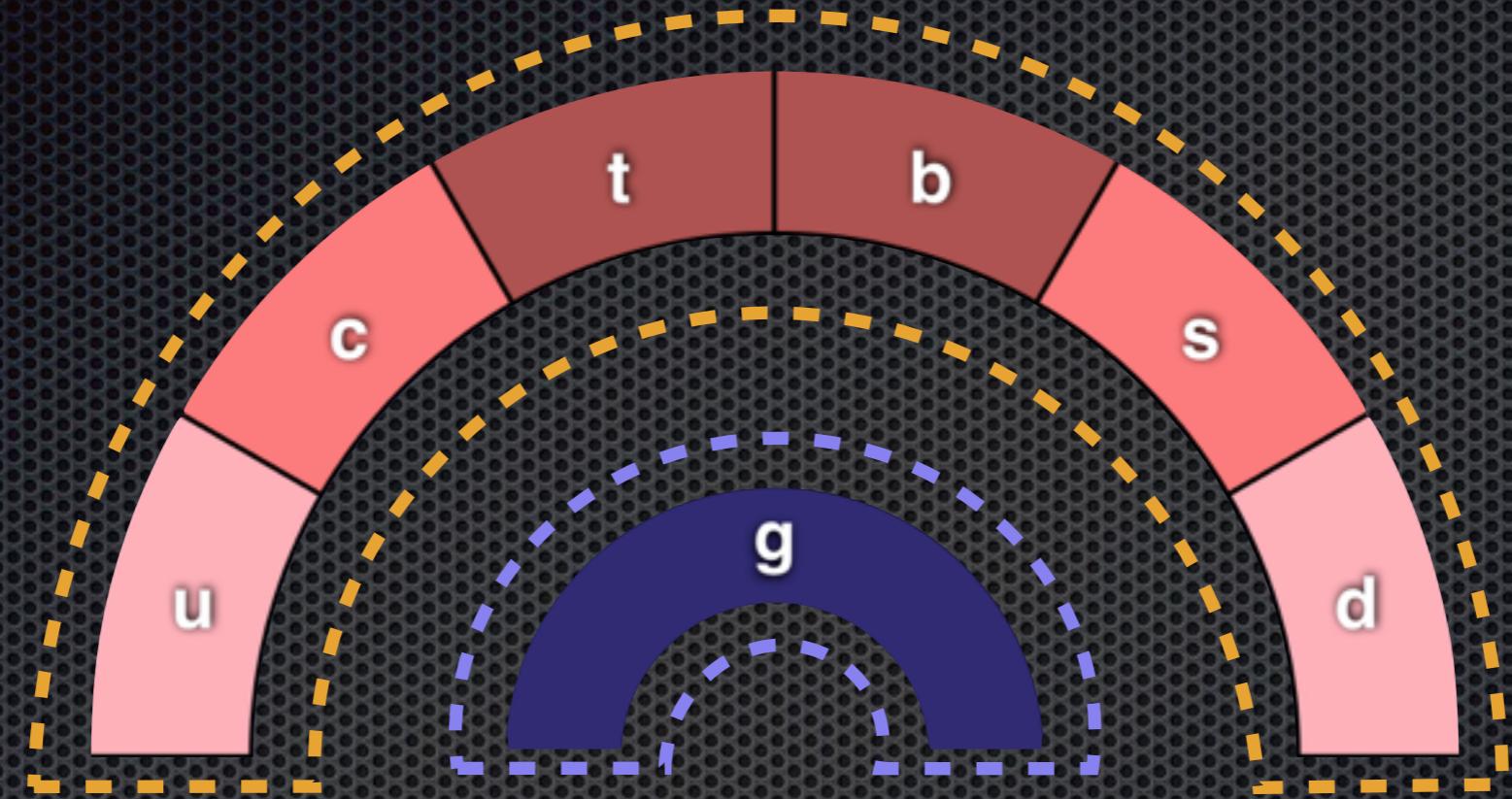


Standard Model of particle physics



strong force

Standard Model of particle physics



color = r, g, b

gluons
color = (r, g, b) + (r, g, b)



strong force

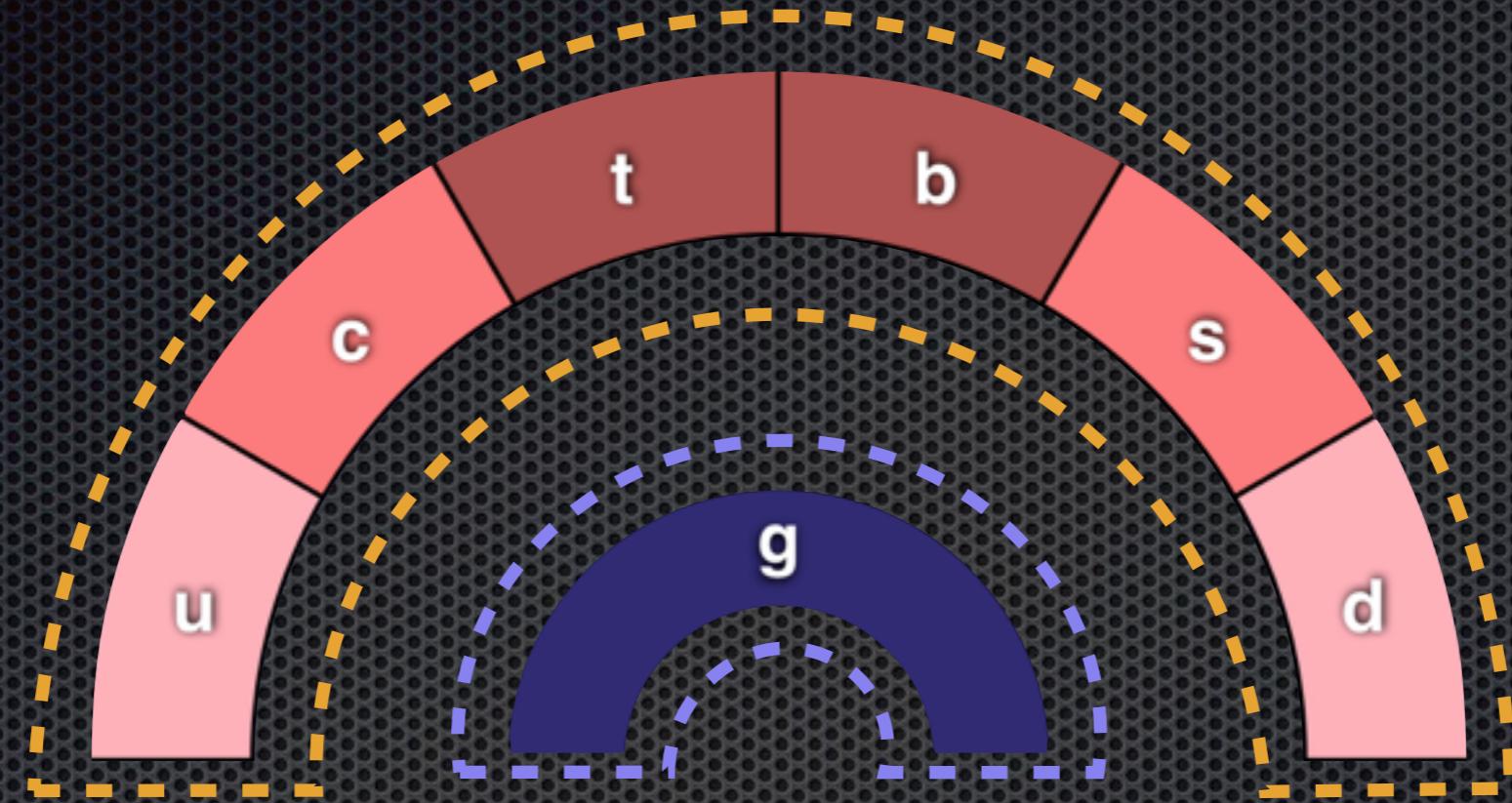


= proton



= neutron

Standard Model of particle physics



color = r, g, b

gluons
color = (r, g, b) + (r, g, b)



strong force



= proton

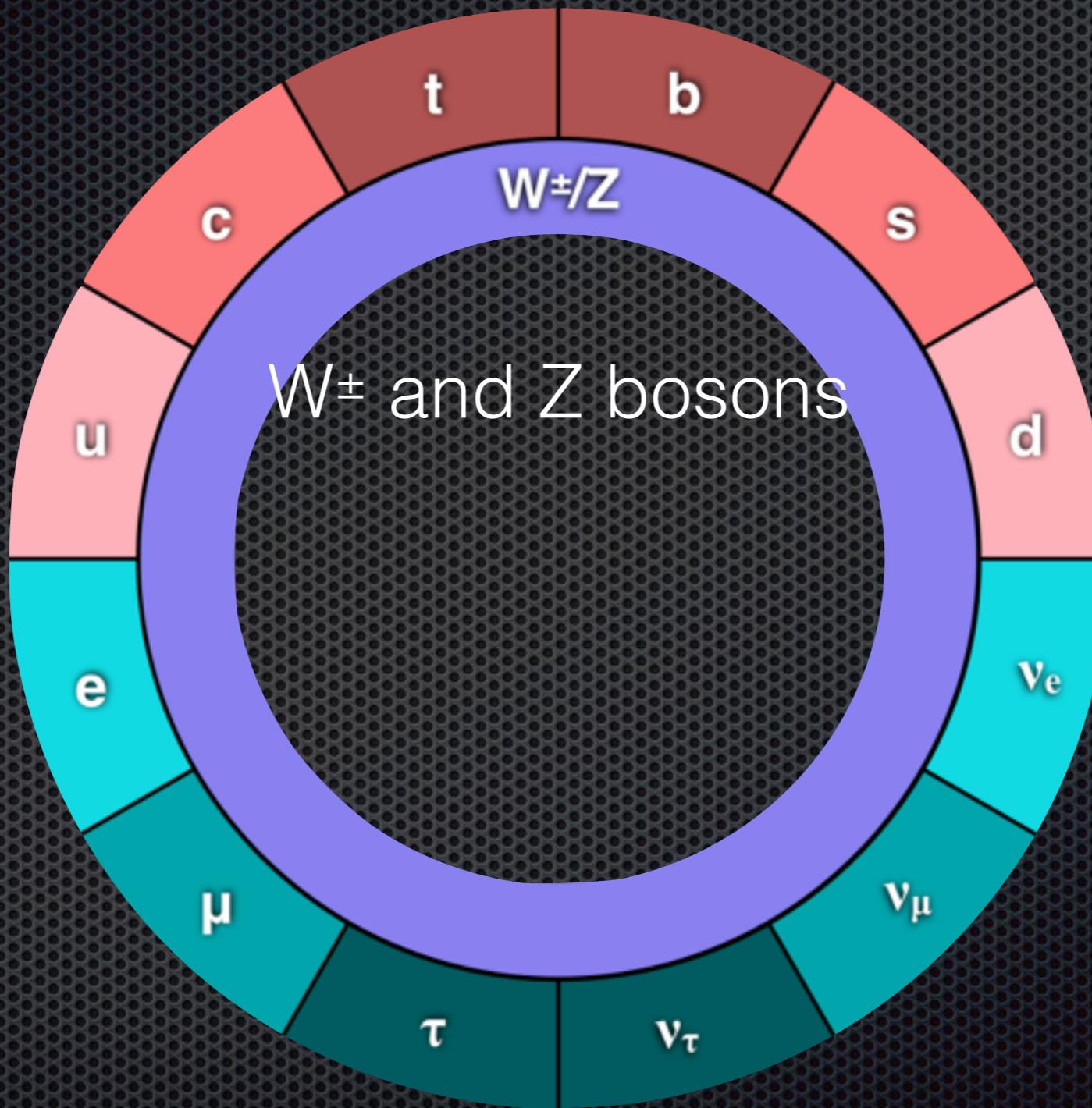


= neutron



= pion

Standard Model of particle physics



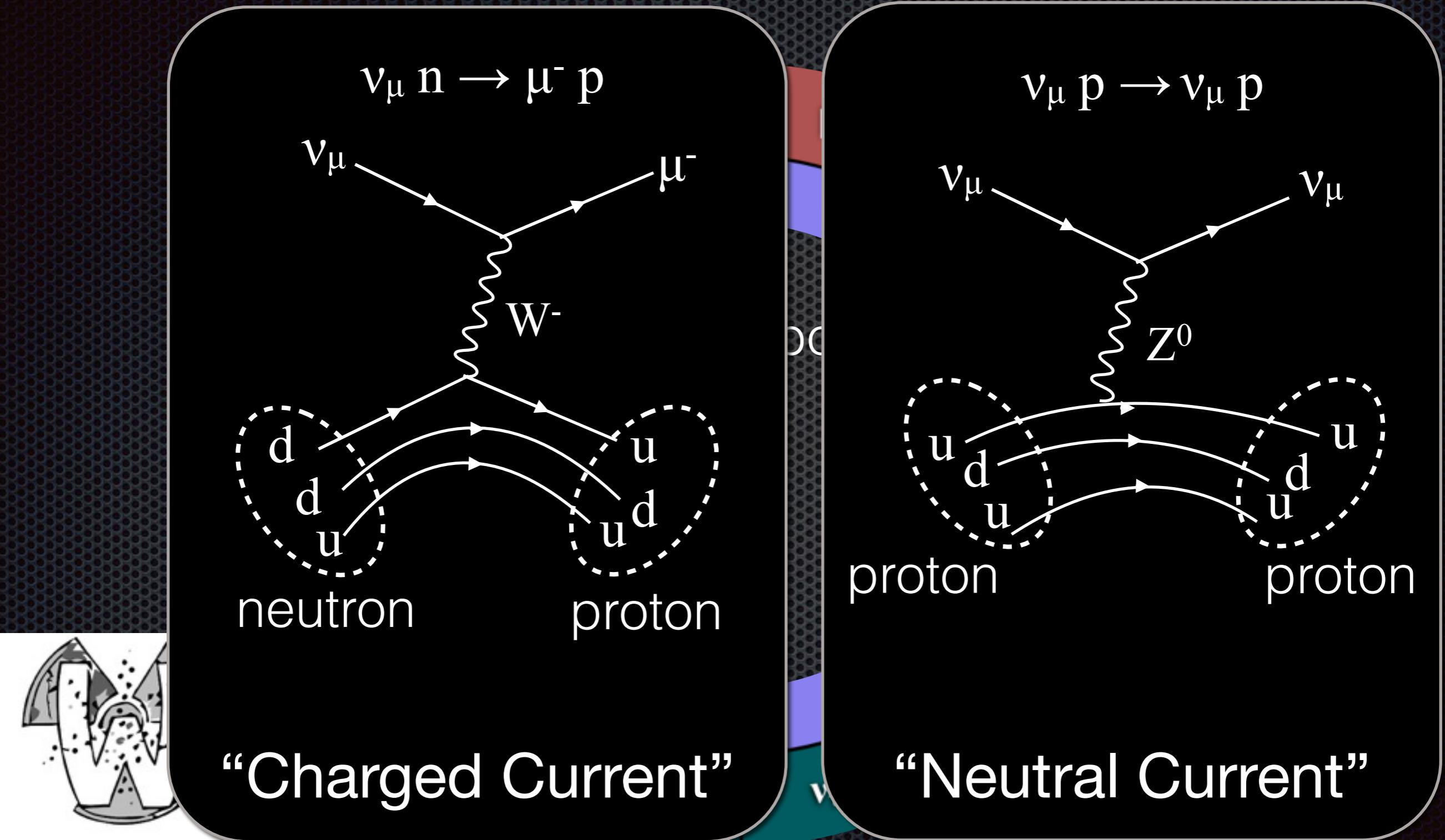
weak force

Standard Model of particle physics



weak force

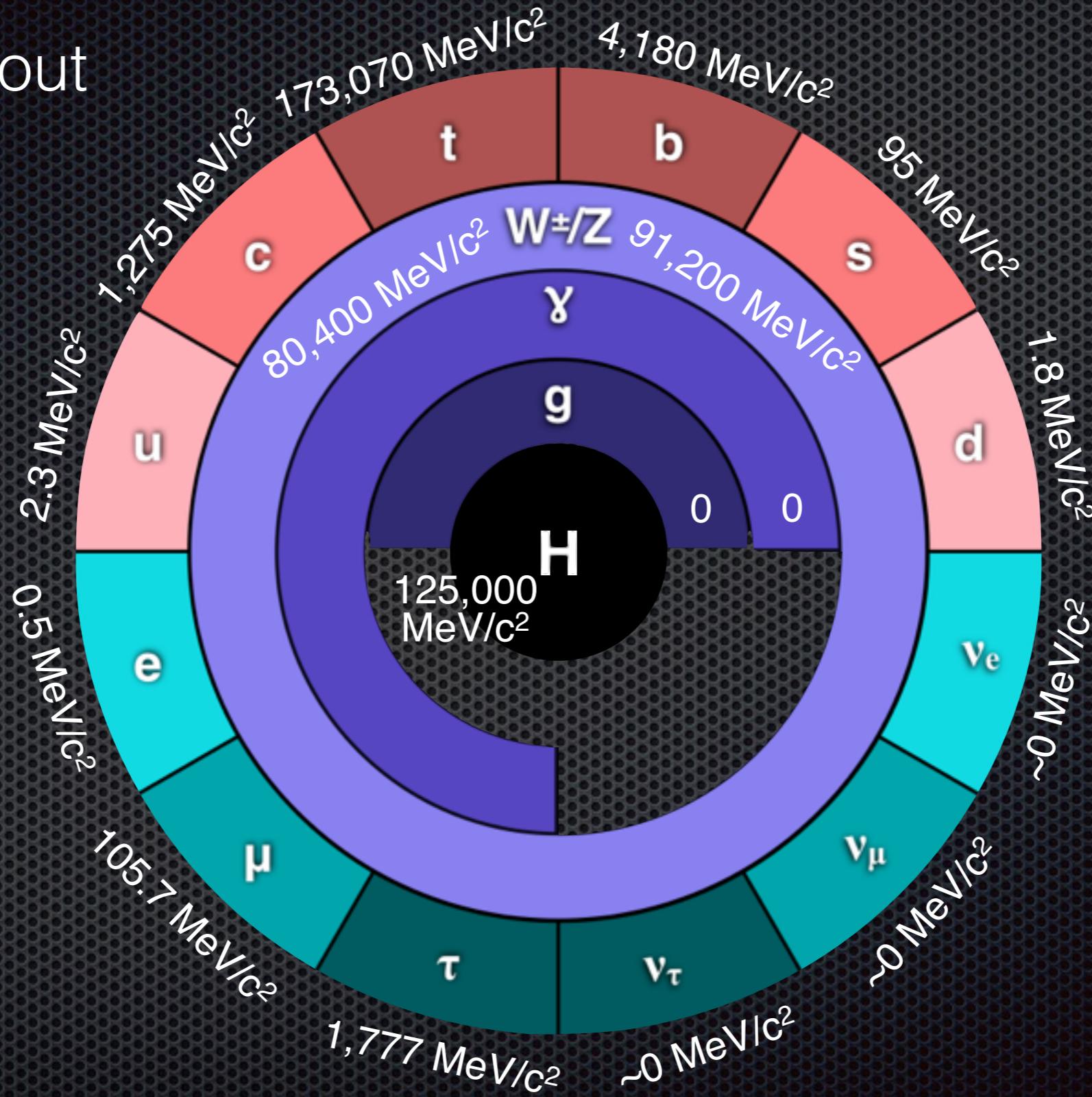
Tia Miceli



weak hypercharge governs

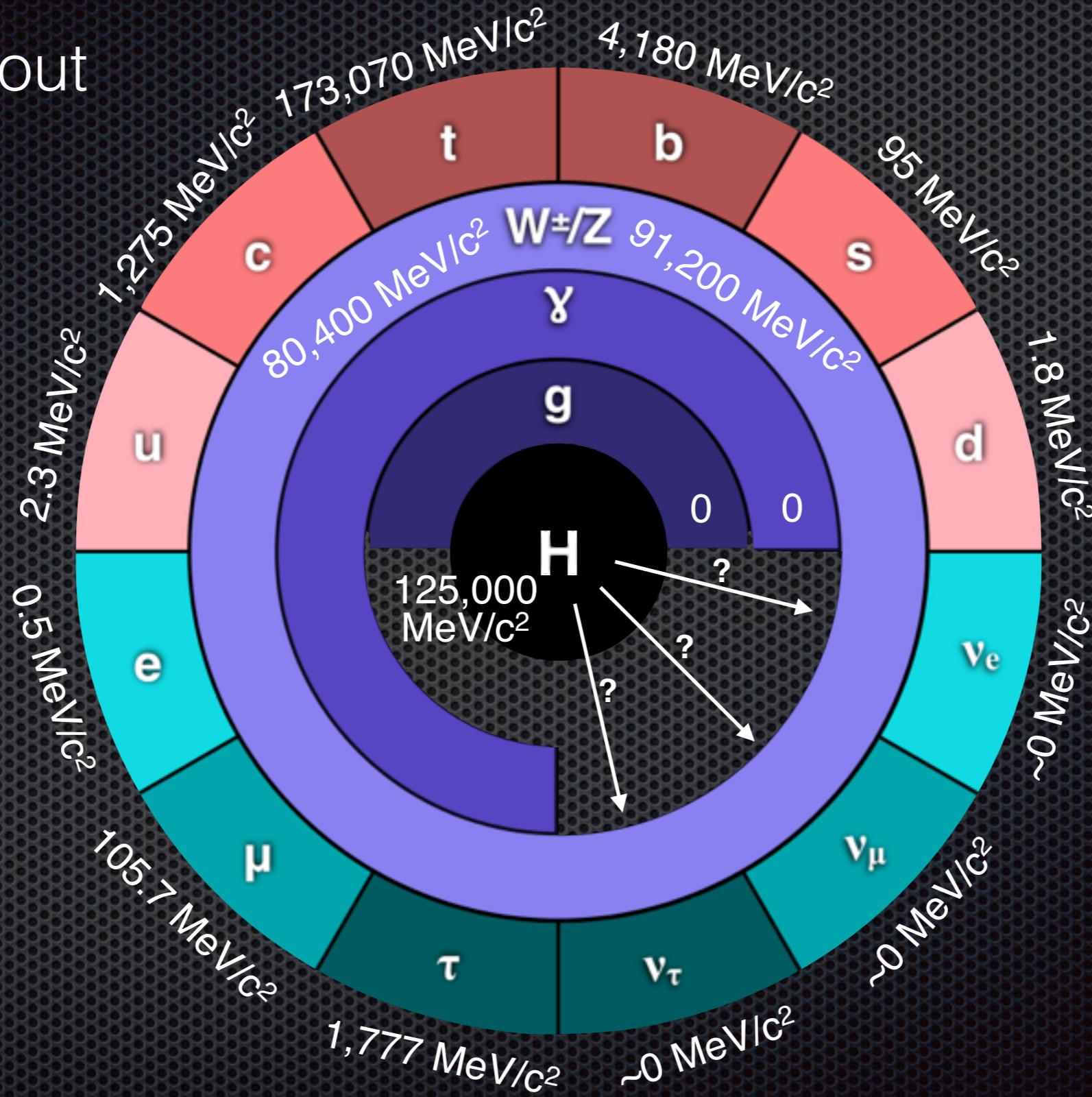
Standard Model of particle physics

A word about mass...



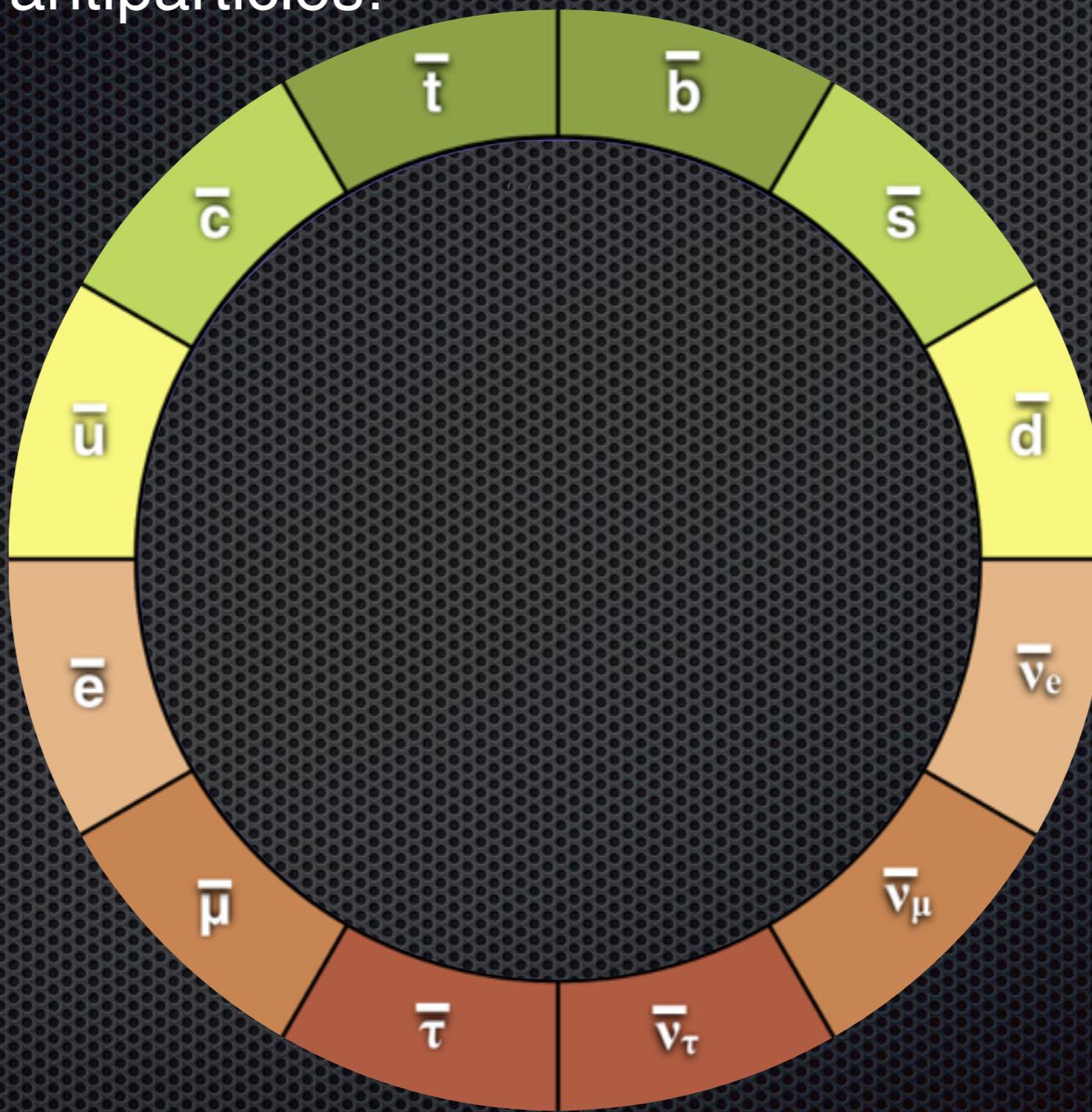
Standard Model of particle physics

A word about mass...



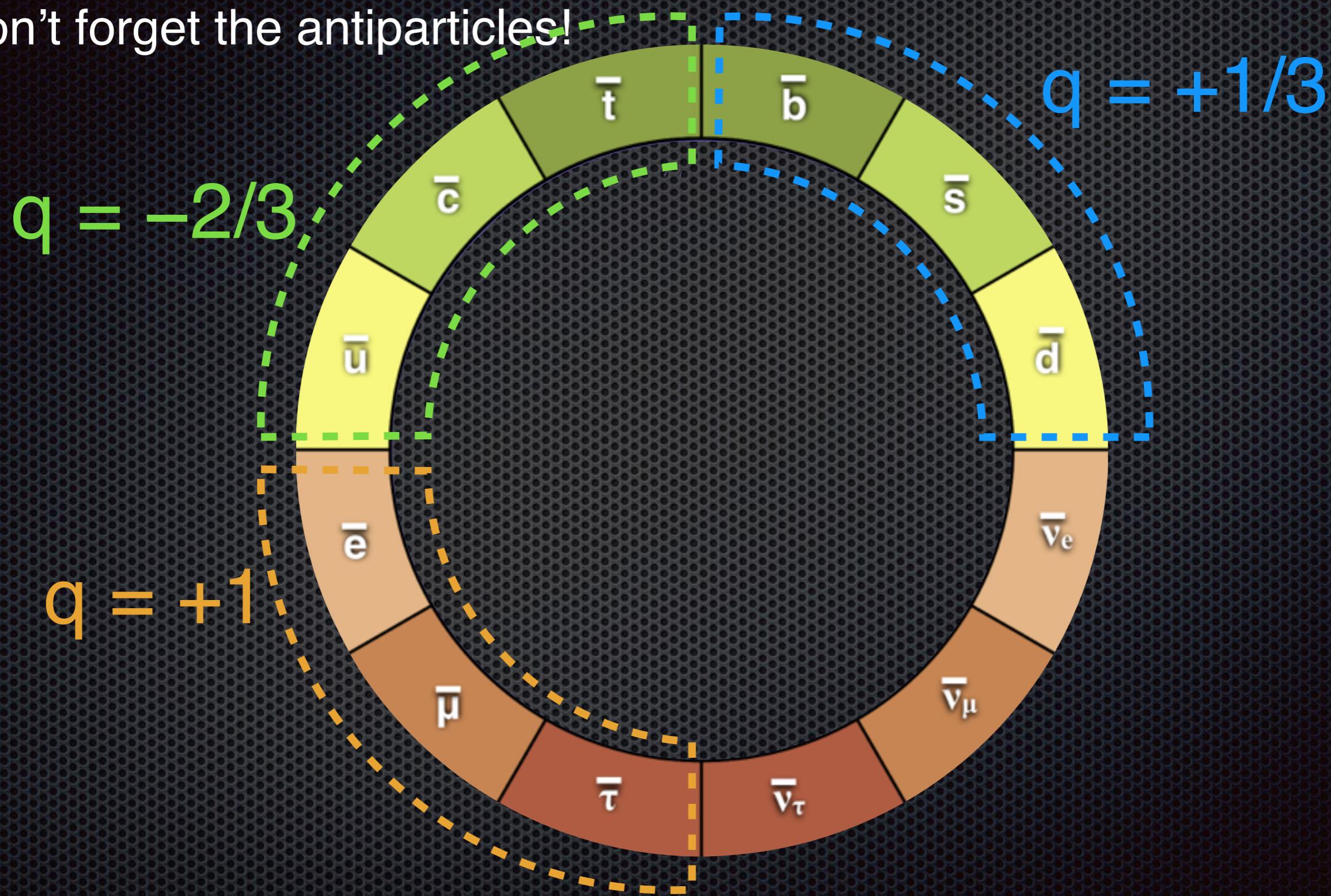
Standard Model of particle physics

Don't forget the antiparticles!



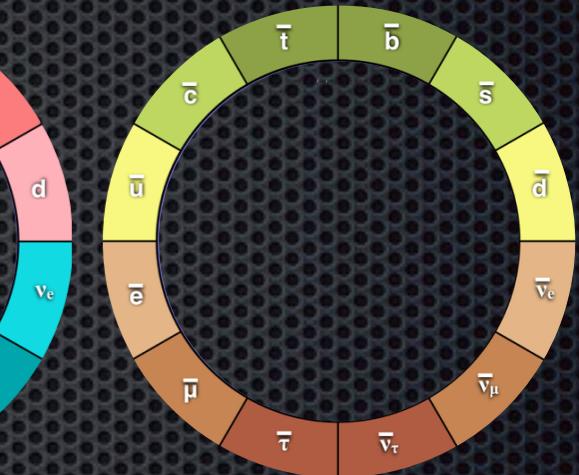
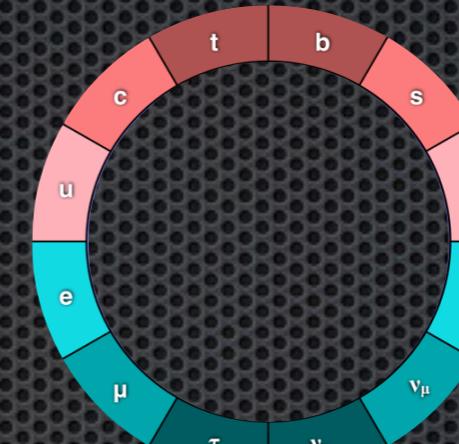
Standard Model of particle physics

Don't forget the antiparticles!



Standard Model of particle physics

- These discovered fermions →



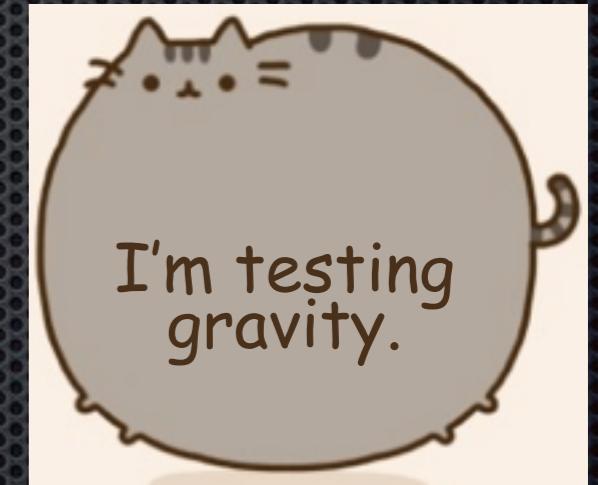
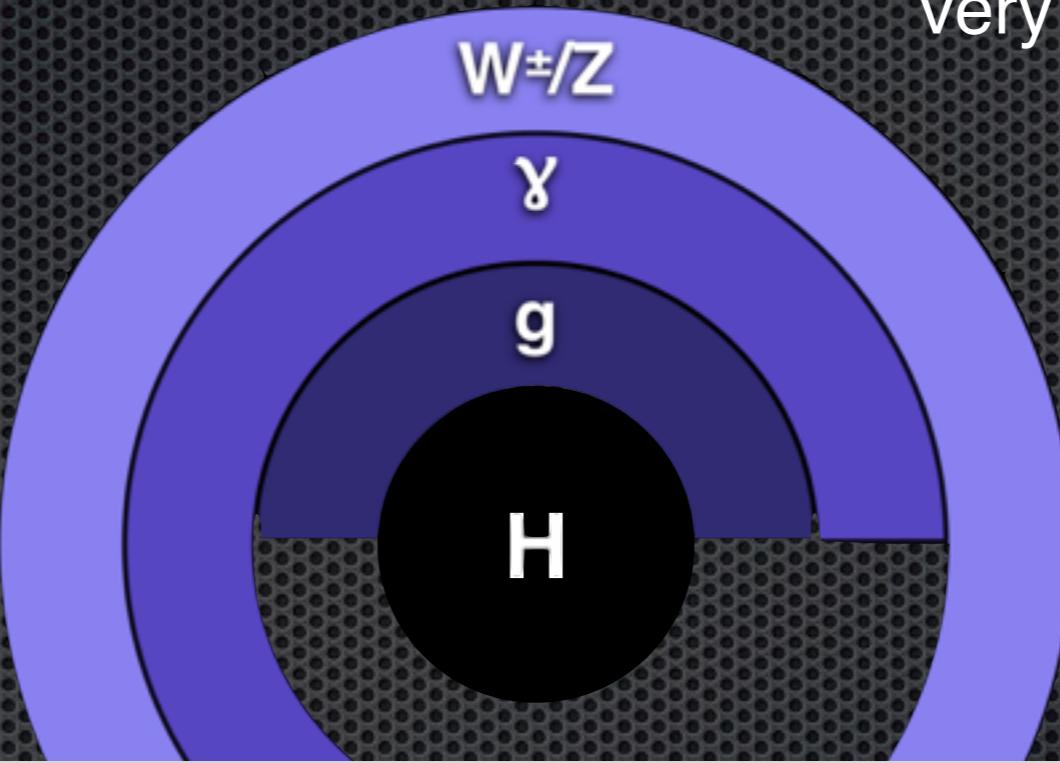
- Whose behavior are described by interactions with these discovered bosons (“force carriers”) using quantum field theory.



Standard Model of particle physics

A word about strength of forces...

P.S. Gravity is not included in the Standard Model, but is very weak. Strength $\sim 10^{-40}$.



force	strength (low E)	distance
strong	1	$10^{-15} \text{ m (nucleus size)}$
e&m	$\frac{1}{137}$	∞
weak	10^{-6}	$10^{-18} \text{ m (0.1% nucleus)}$

Neutrinos are hard to detect!

the weak force is **weak**!
neutrinos interact
100,000,000,000
times less often than quarks



A neutrino has a good chance of traveling through
200 earths before interacting at all

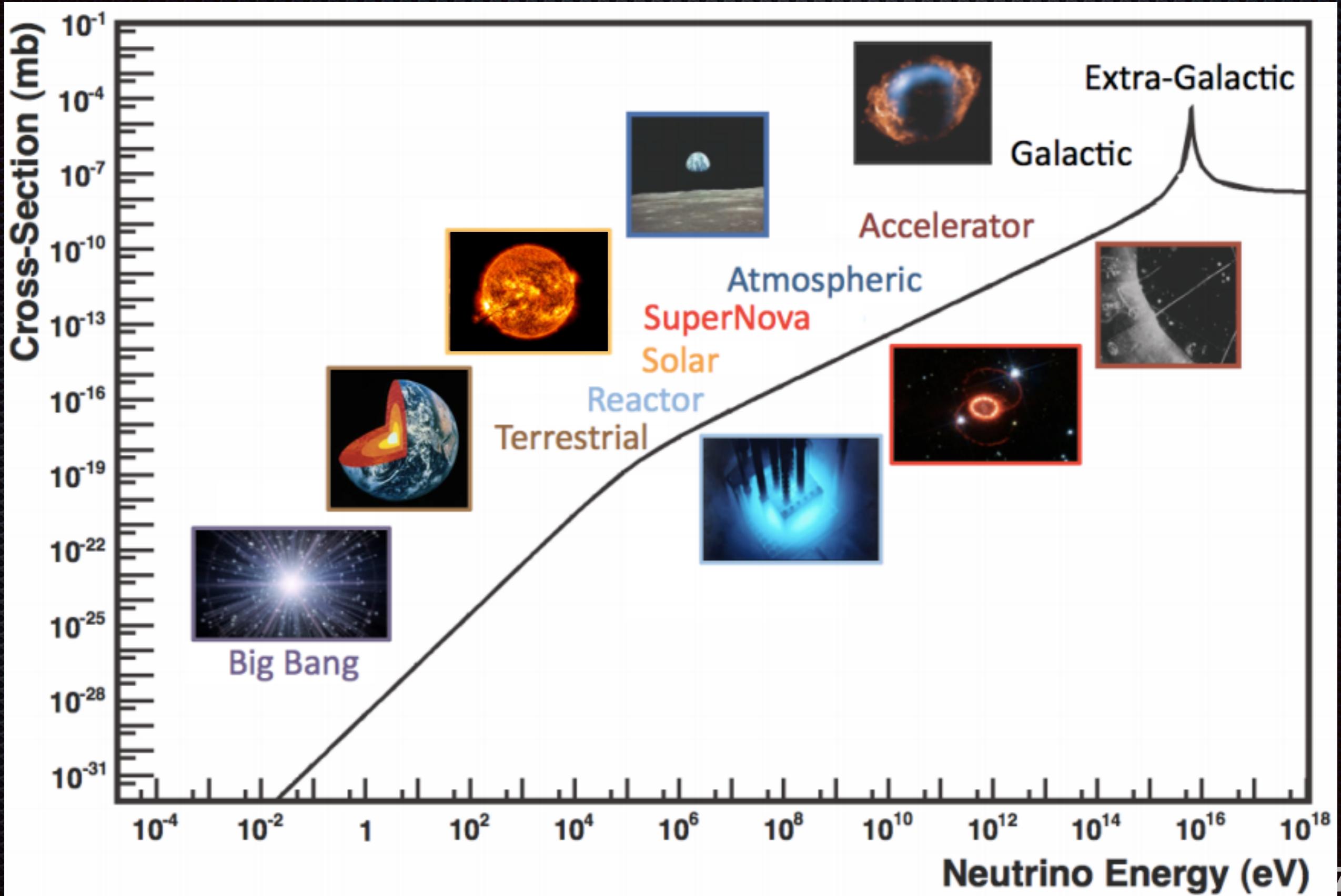
Because they interact rarely...

- We increase our chances of observation by studying very prolific neutrino sources with gigantic detectors!



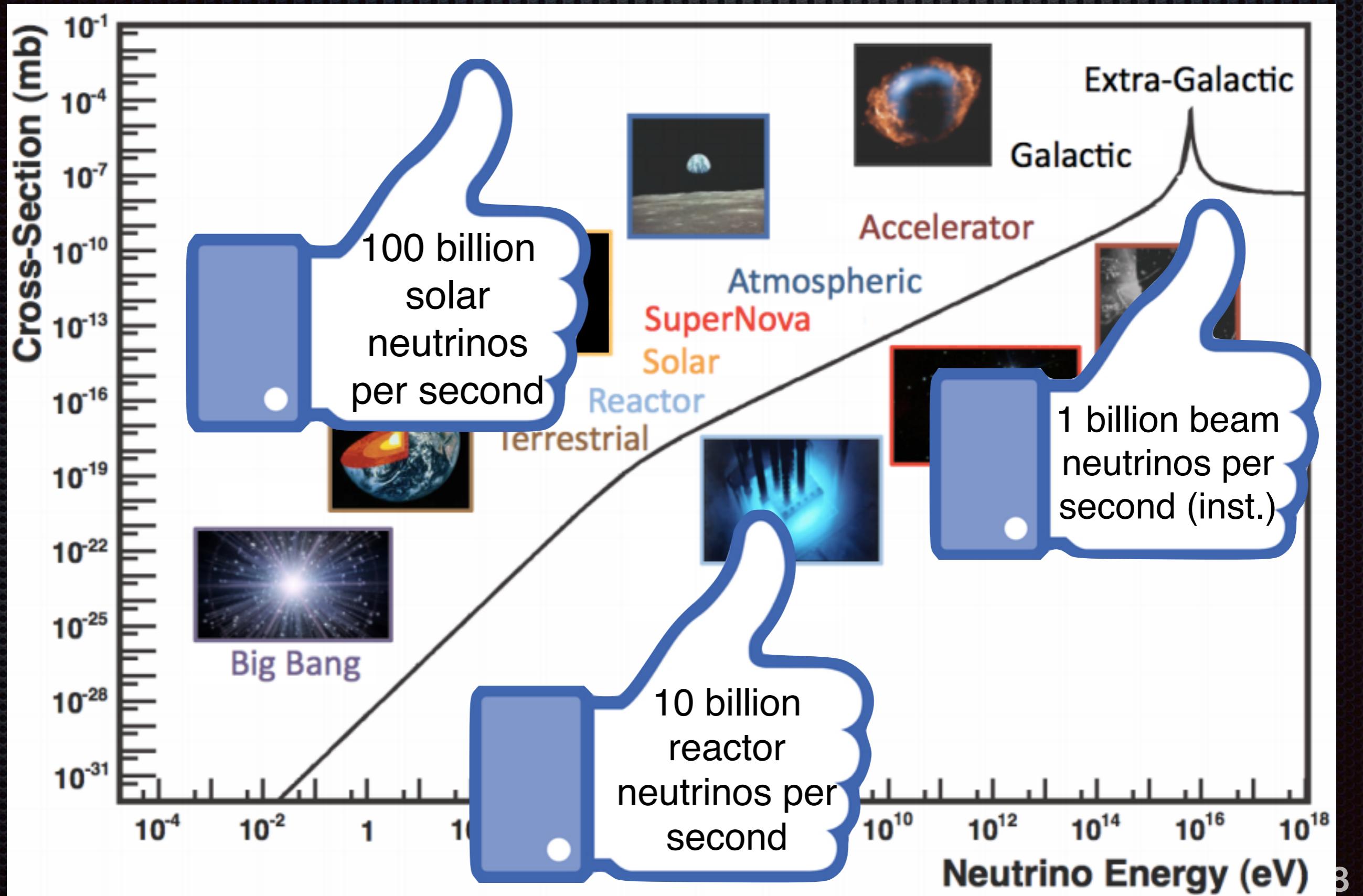
Neutrino sources

Probability of Interaction

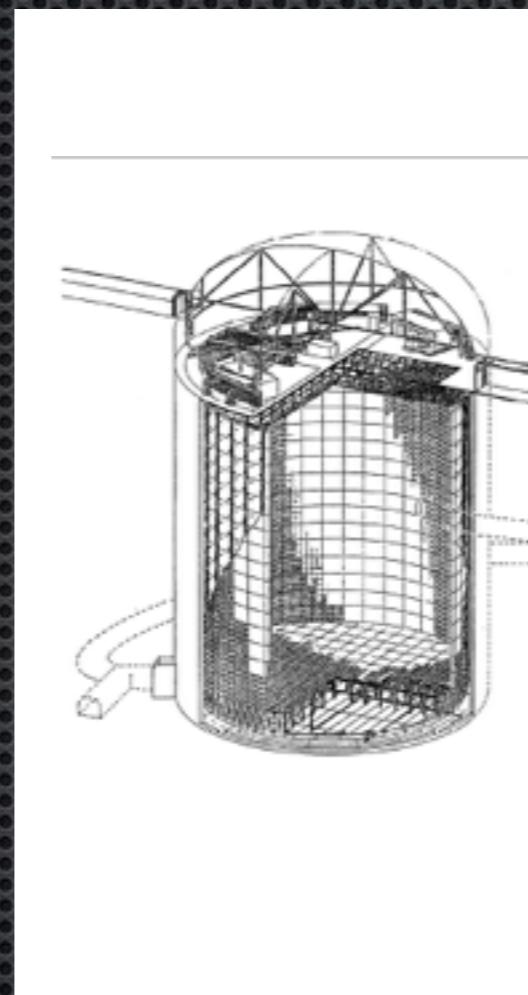
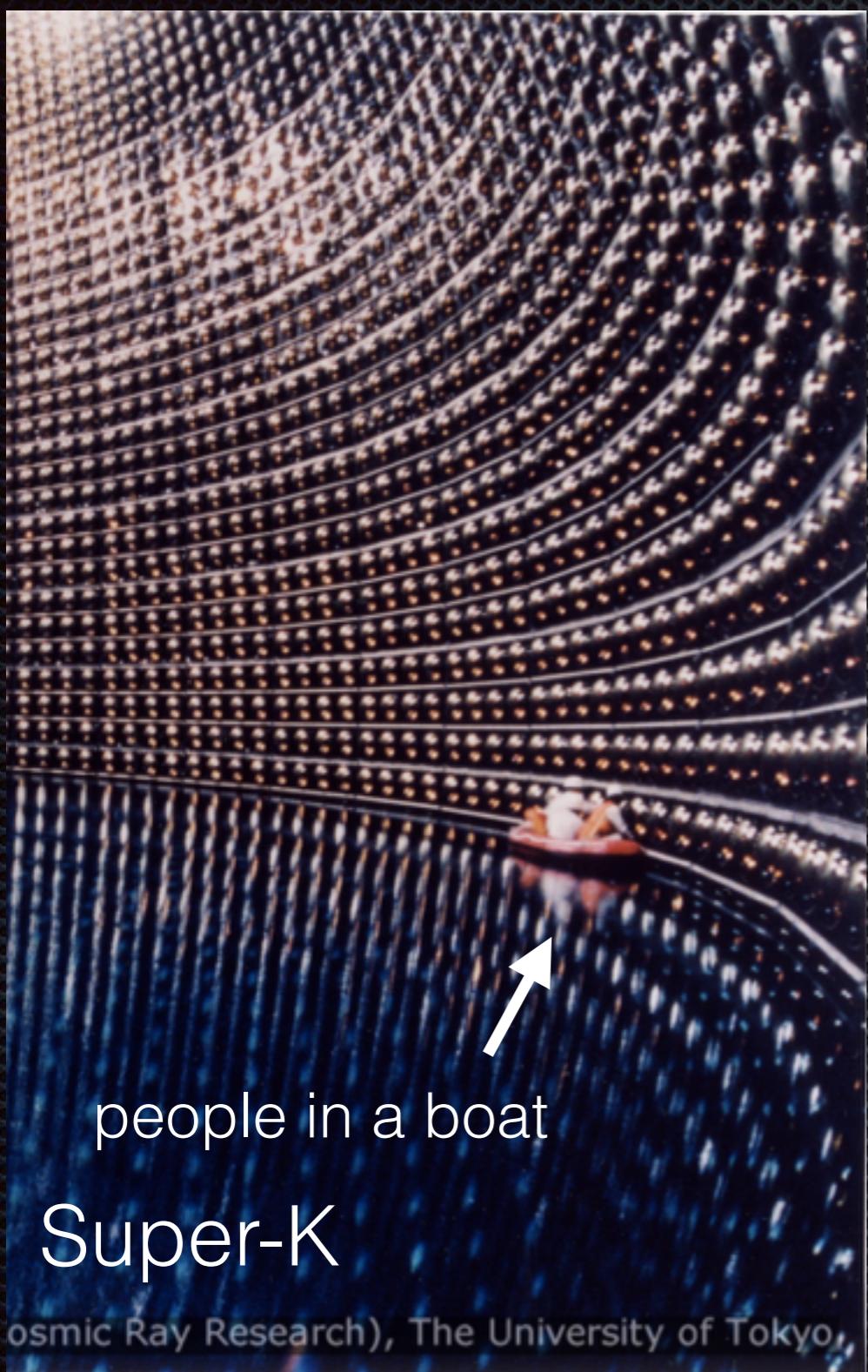


Neutrino sources

Probability of Interaction

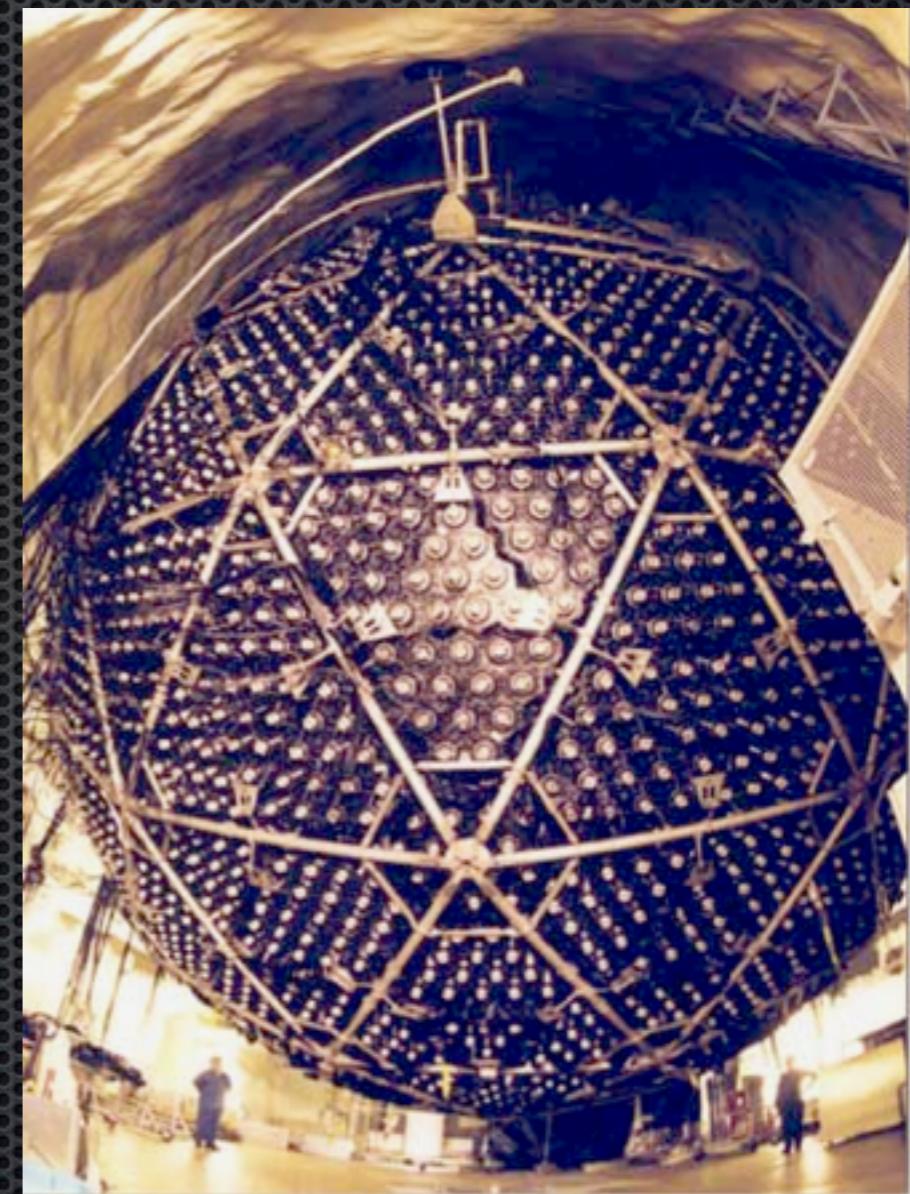


Scale of neutrino experiments



Solar

Scale of neutrino experiments



Aspen East

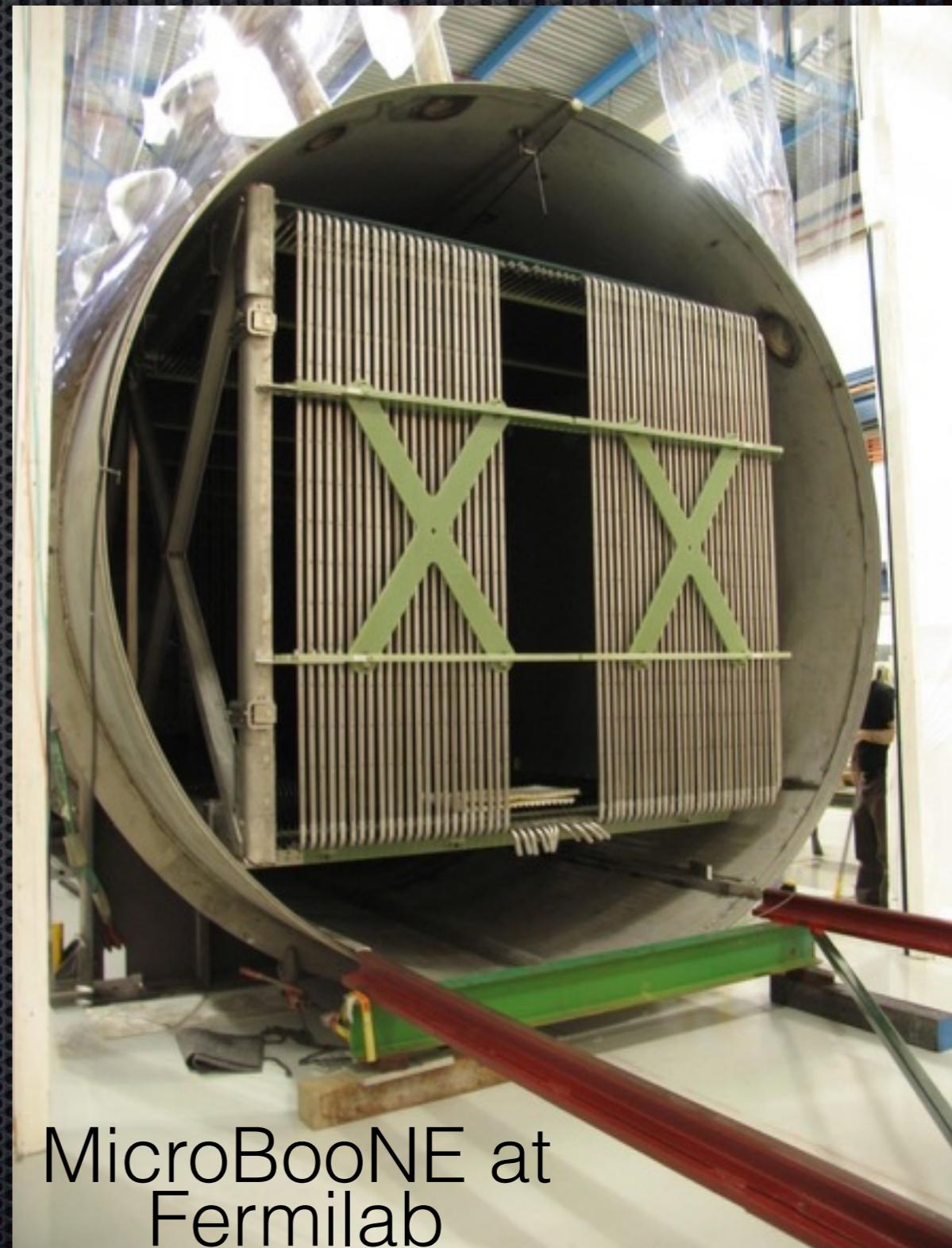
Solar

Sudbury Neutrino Experiment
(SNO)

Scale of neutrino experiments



Beam

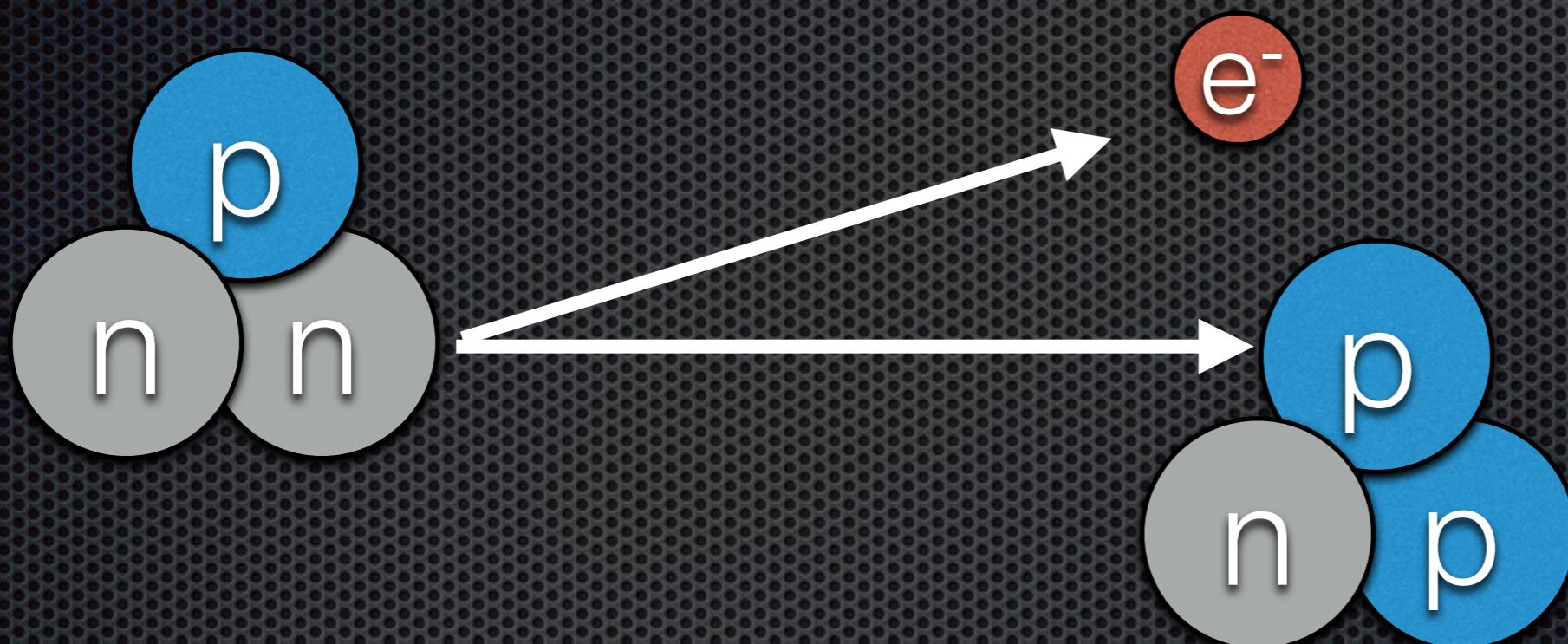


MicroBooNE at
Fermilab

- What's particle physics? (translated into “Tia Speak”)
 - Where are we in neutrino physics? (Story time!)
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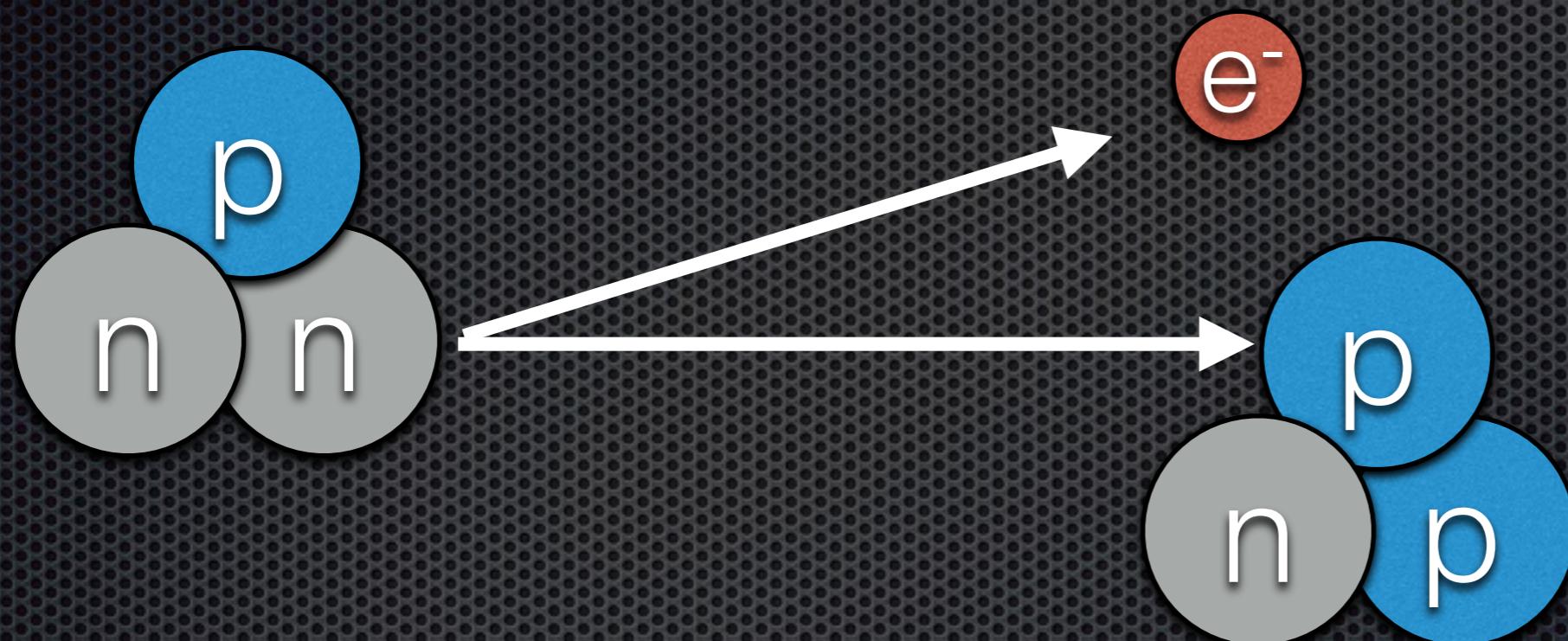
At the beginning, there was beta decay...

Ex. Tritium \rightarrow Helium + e⁻



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Ex. Tritium \rightarrow Helium + e⁻

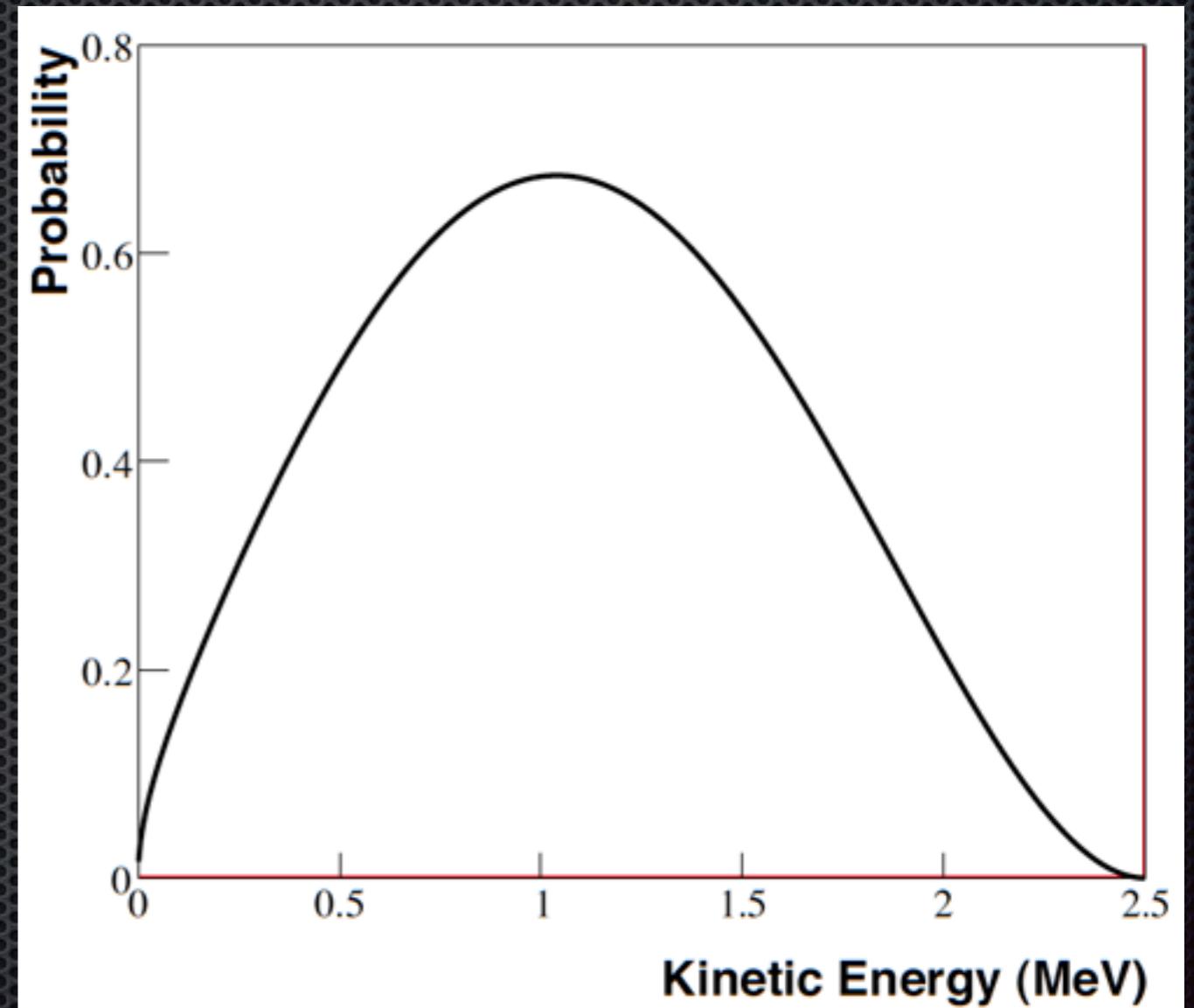


$$E_{\text{initial}} > E_{\text{final}}$$

products were missing energy

Observation of e^- energy spectrum

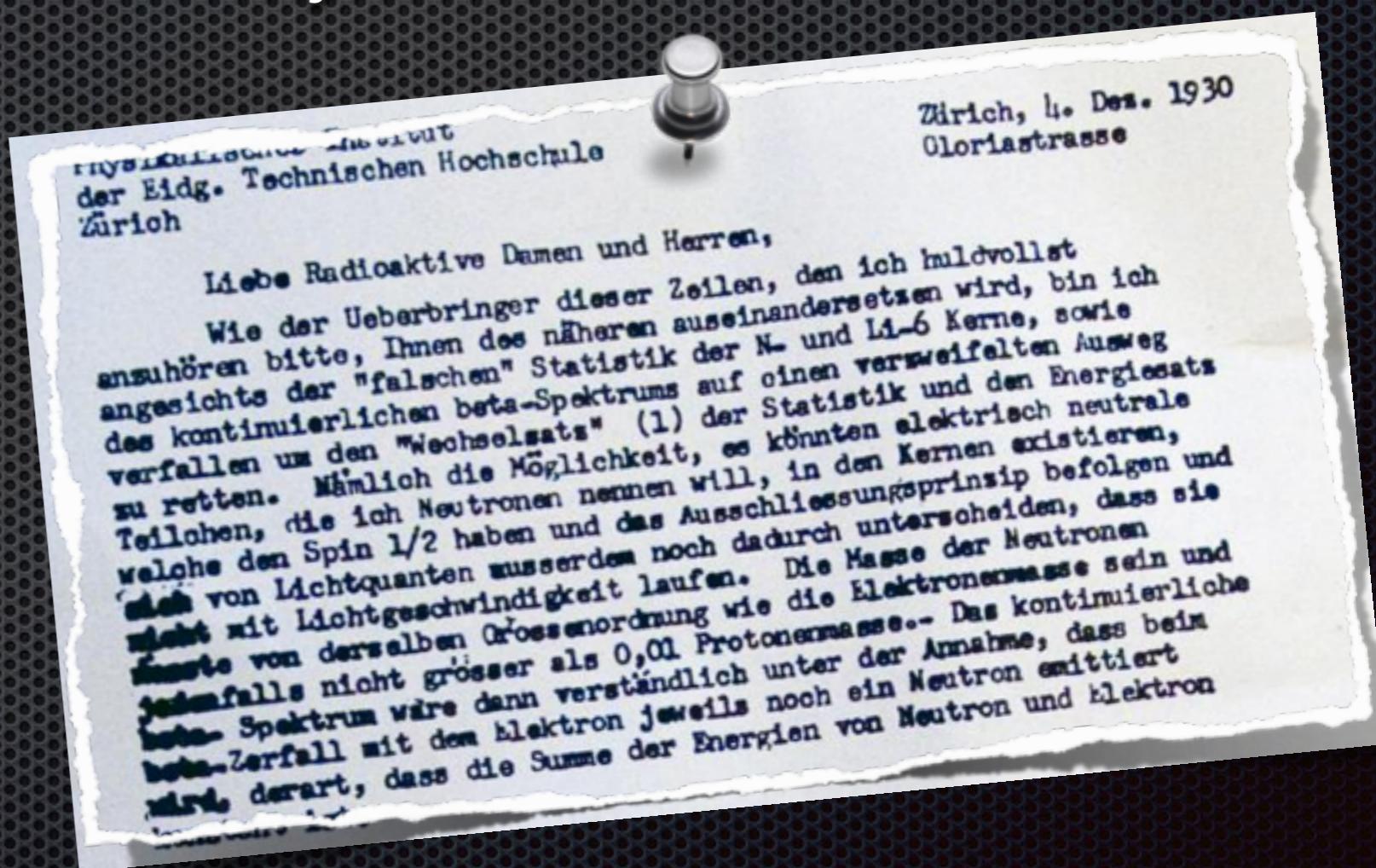
- Energy conservation was either violated, or a hidden particle carried away energy.



Sir James Chadwick 1914

A desperate remedy!

- Energy and spin 1/2 carried away by a light “invisible” particle.
- Later named “neutrino” by Enrico Fermi



Wolfgang Pauli, 1930

Letter to Zurich Conference

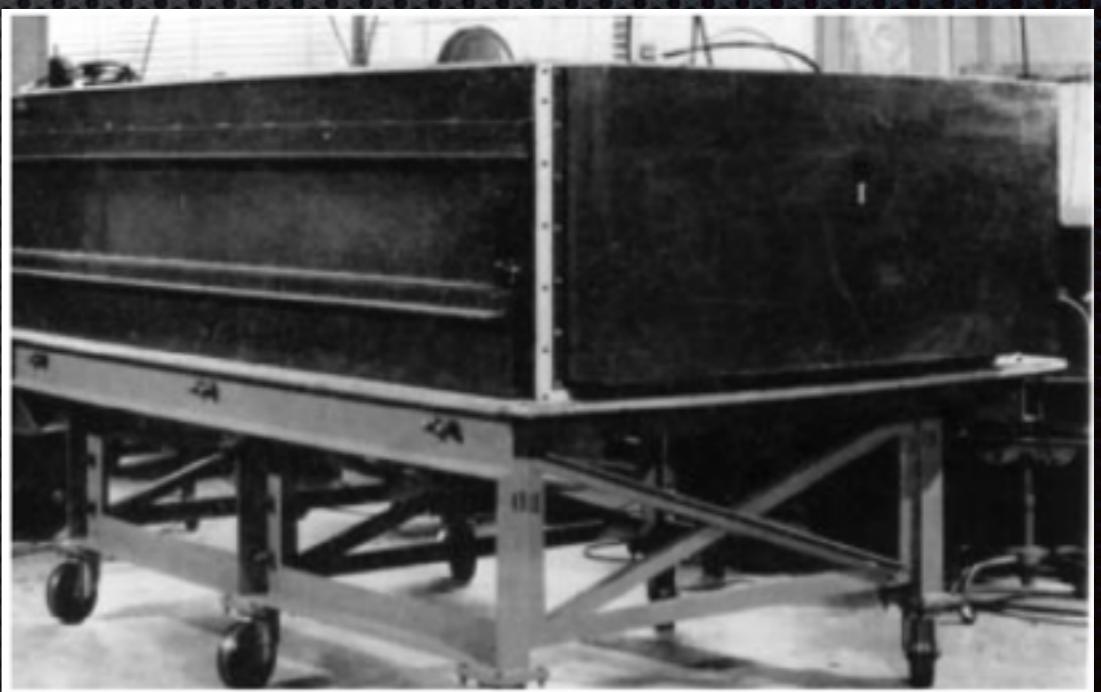
Neutrino Wanted!

- “Project Poltergeist”



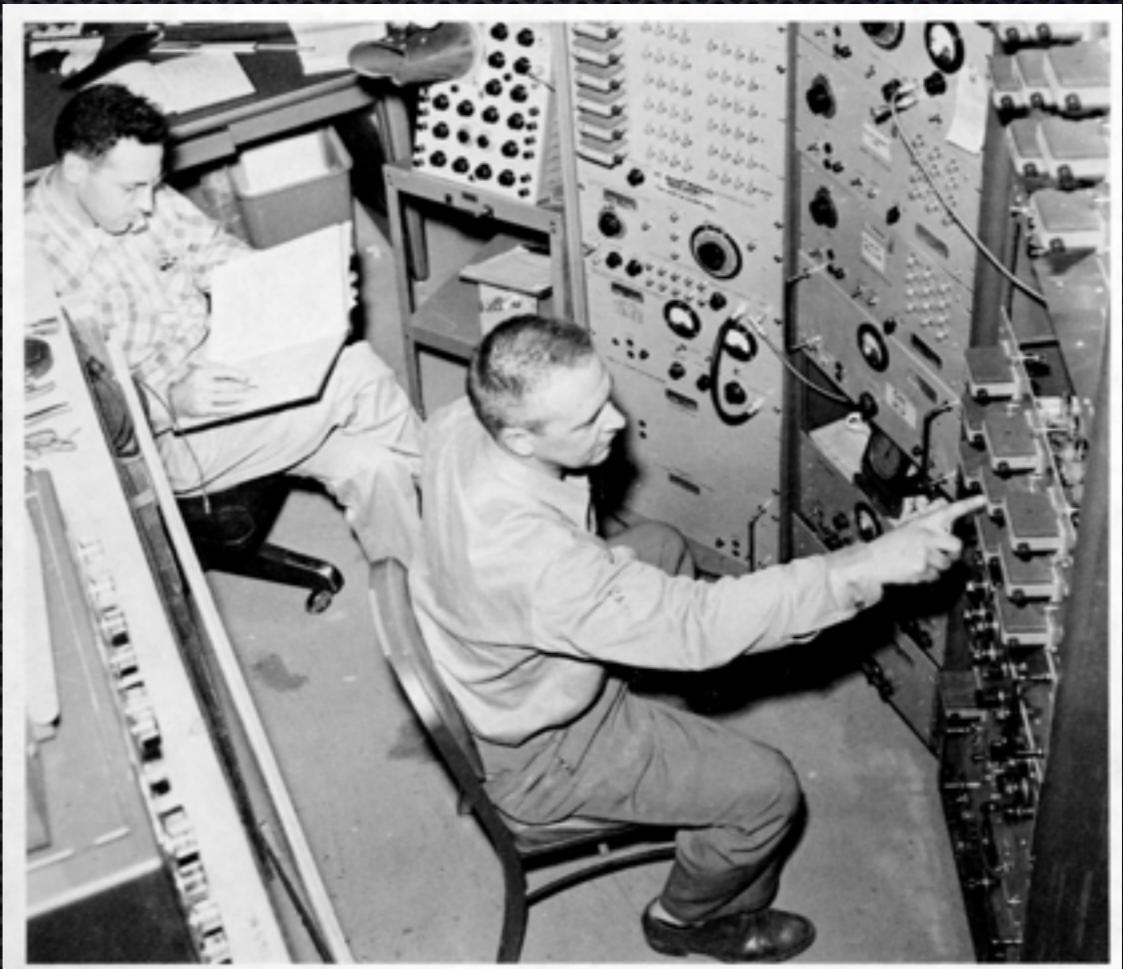
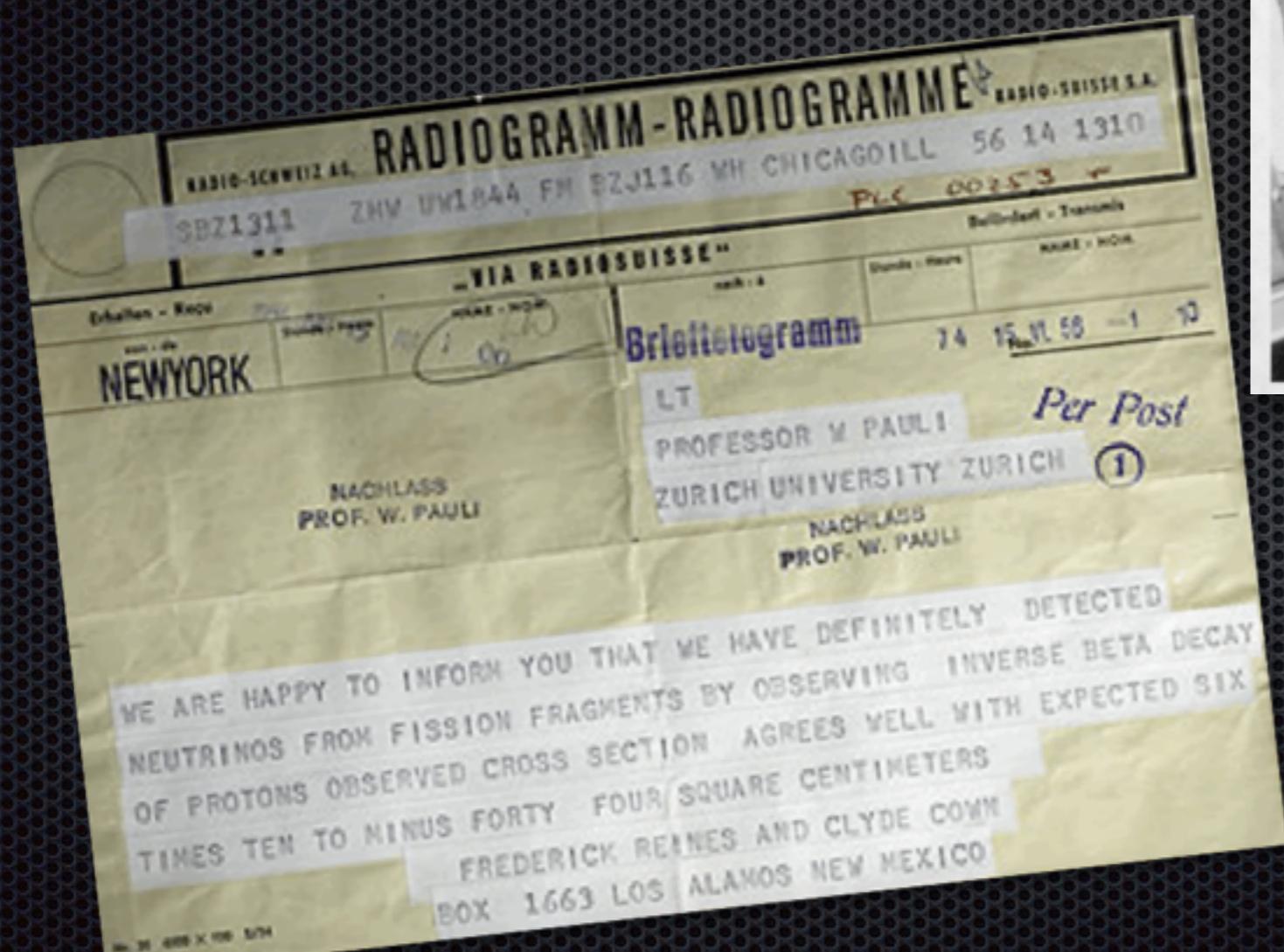
Savannah River
Nuclear Reactor,
South Carolina. 1955

Neutrino Detector:
tank of water and
liquid scintillator



Neutrino Caught!

- $\bar{\nu}_e$ discovered 14 June 1956



Clyde Cowan &
Frederick Reines



awarded 1995

to Frederick for detection
of the neutrino

Spark chamber reveals ν_μ

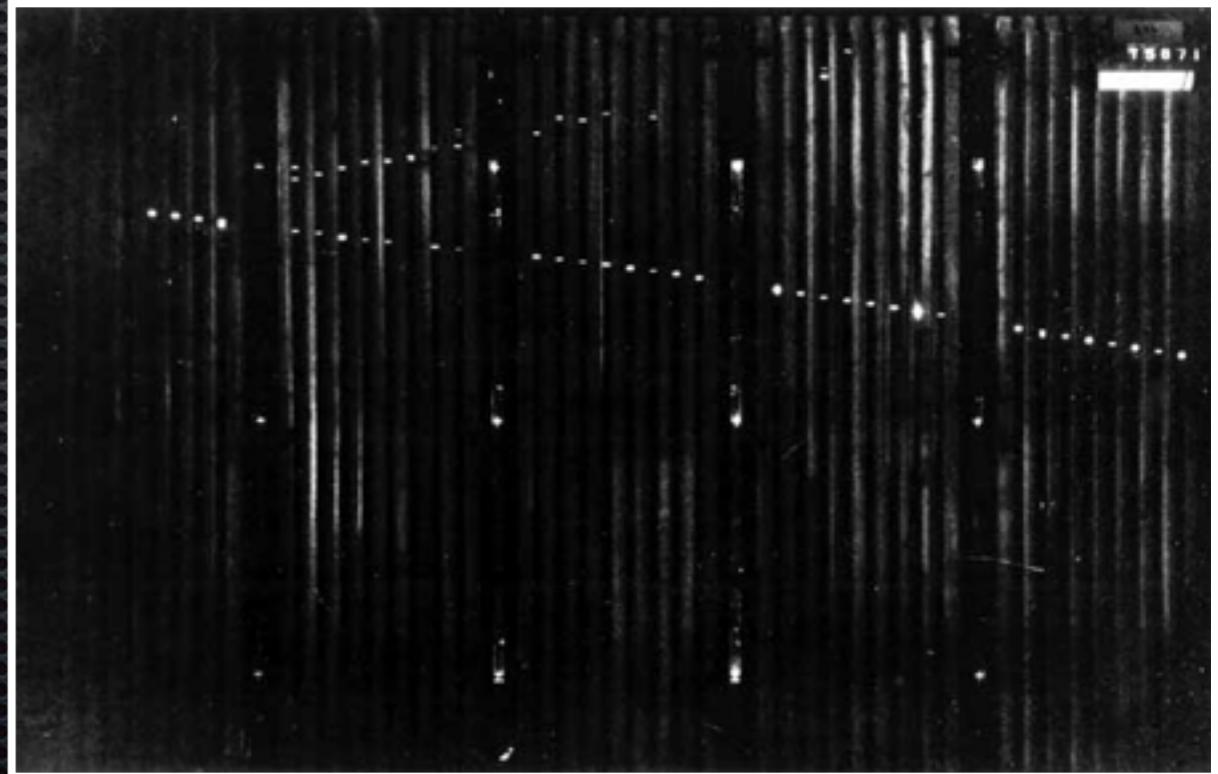
- ν_μ suspected, then discovered in 1962



Muons created by a muon neutrino beam in a spark chamber



Leon M. Lederman, Melvin Schwartz and Jack Steinberger





awarded 1998

to Lederman, Schwartz and Steinberger
for discovery of the ν_μ

Solar fusion should emit ν_e

- Bahcall, under current solar models, calculated the number of electron flavor neutrinos coming from sun.



John Bahcall

Teaming up!

- 1968, South Dakota.
- Only saw ~1/3 of the expected rate!



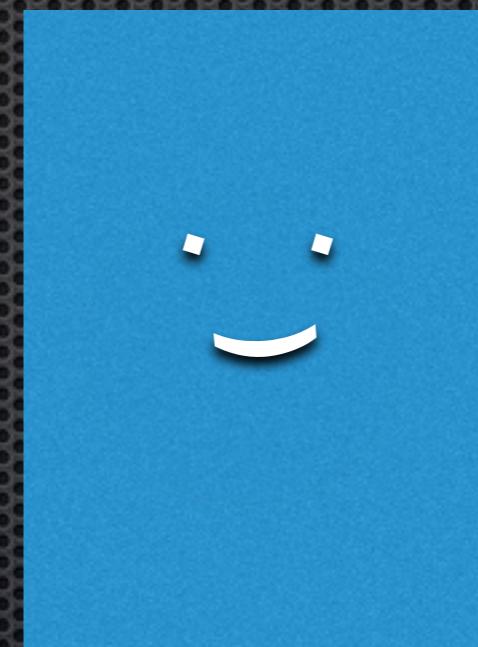
Homestake Experiment, 1960s



Ray Davis and John Bahcall

Neutrino mixing postulated

- Just a few years earlier, such a deficit was predicted due to different types of neutrinos mixing. (PMNS matrix!)
- Possibly the missing electron neutrinos oscillated to another flavor.



Bruno
Pontecorvo
Tia Miceli

Ziro
Maki

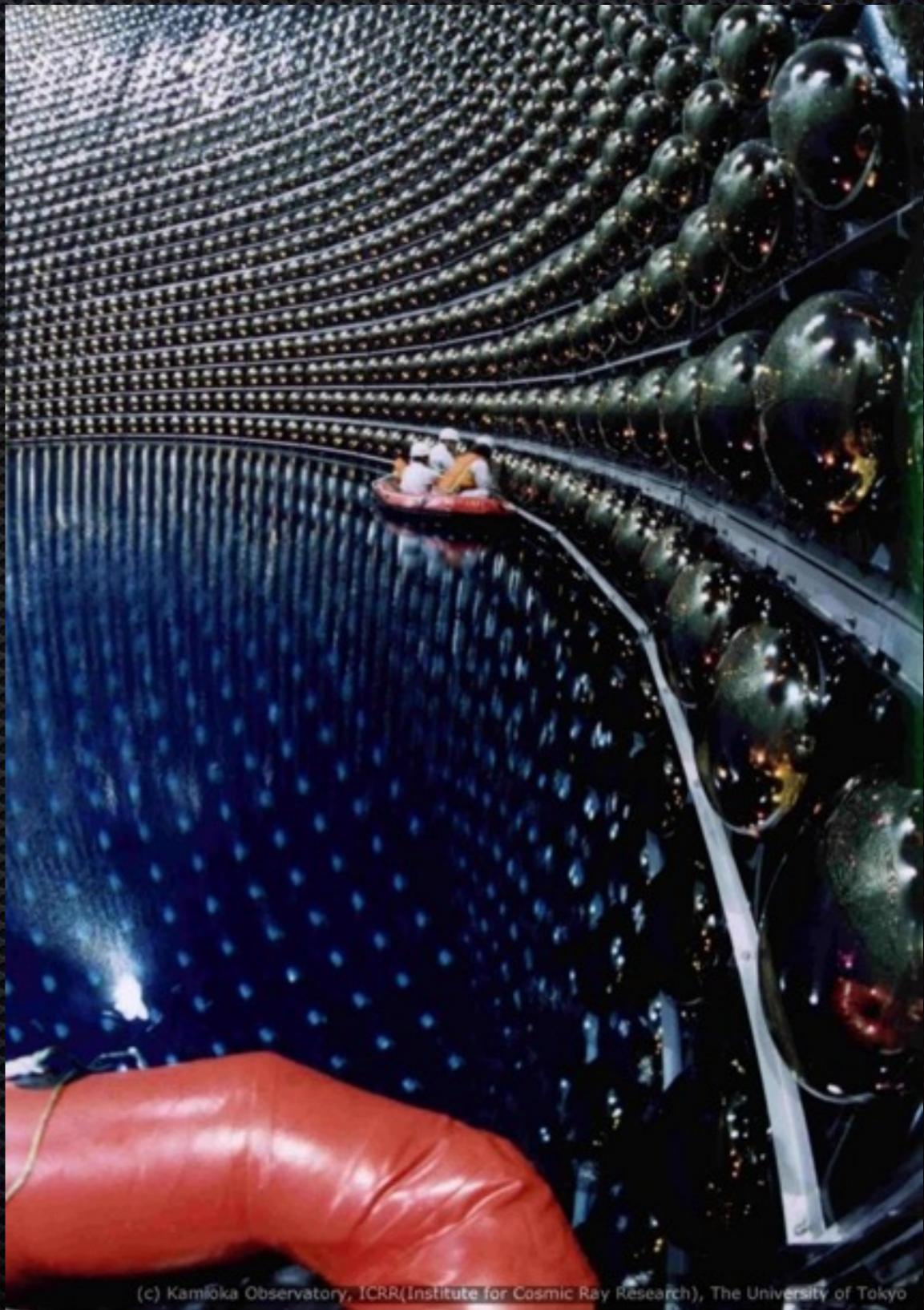
Masami
Nakagawa

Shoichi
Sakata

Conclusive evidence

- 3 July 1998, Japan's Super-K
- Water chernkov detector, sensitive to electron and muon flavor neutrinos.

ν_e sun $\rightarrow \nu_e, \nu_\mu \dots$



(c) Kamioka Observatory, ICRR(Institute for Cosmic Ray Research), The University of Tokyo

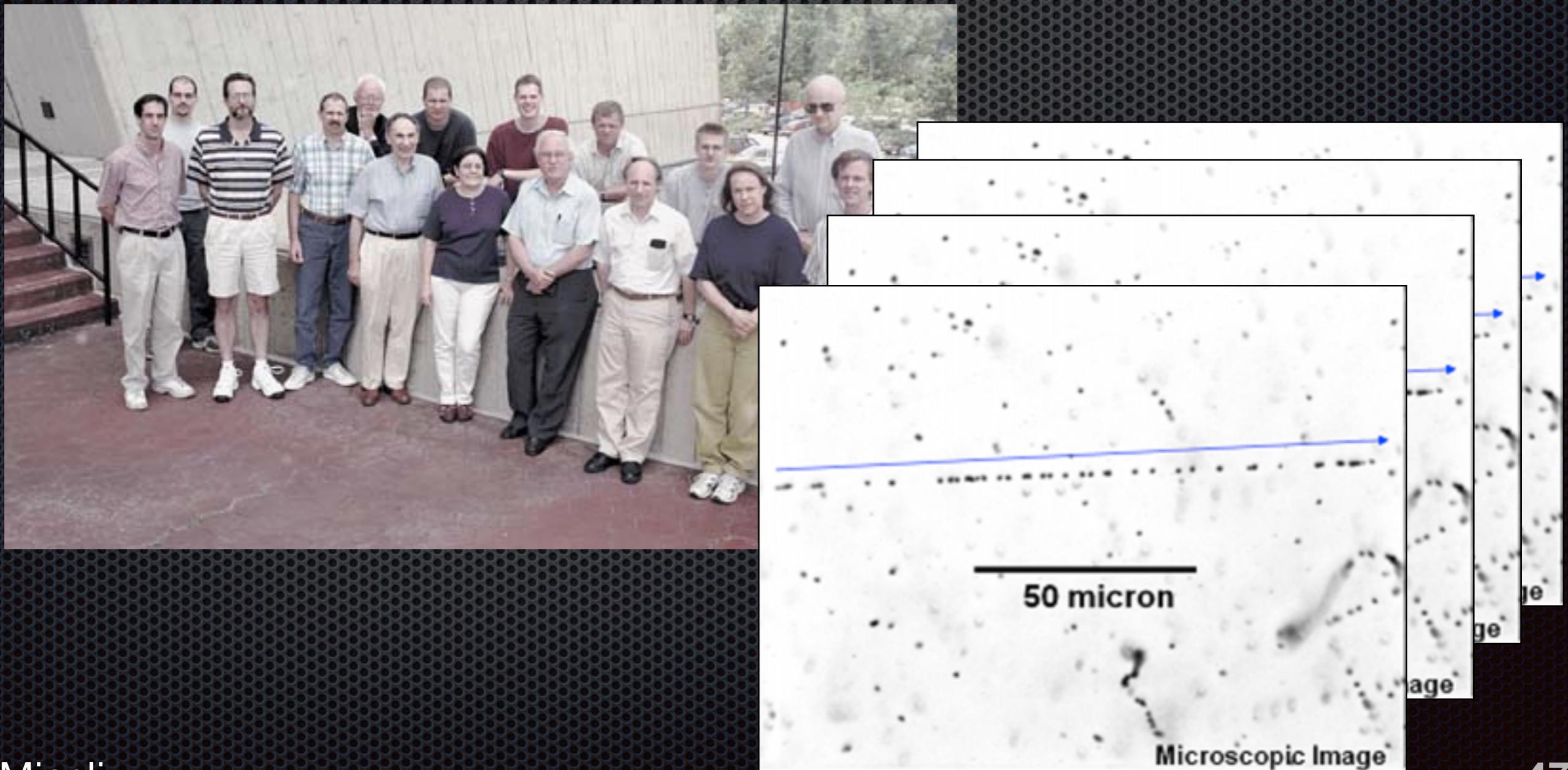


awarded 2002

to Davis and Koshiba
for “cosmic neutrinos” (solar, atmospheric, supernova)

Tau neutrino

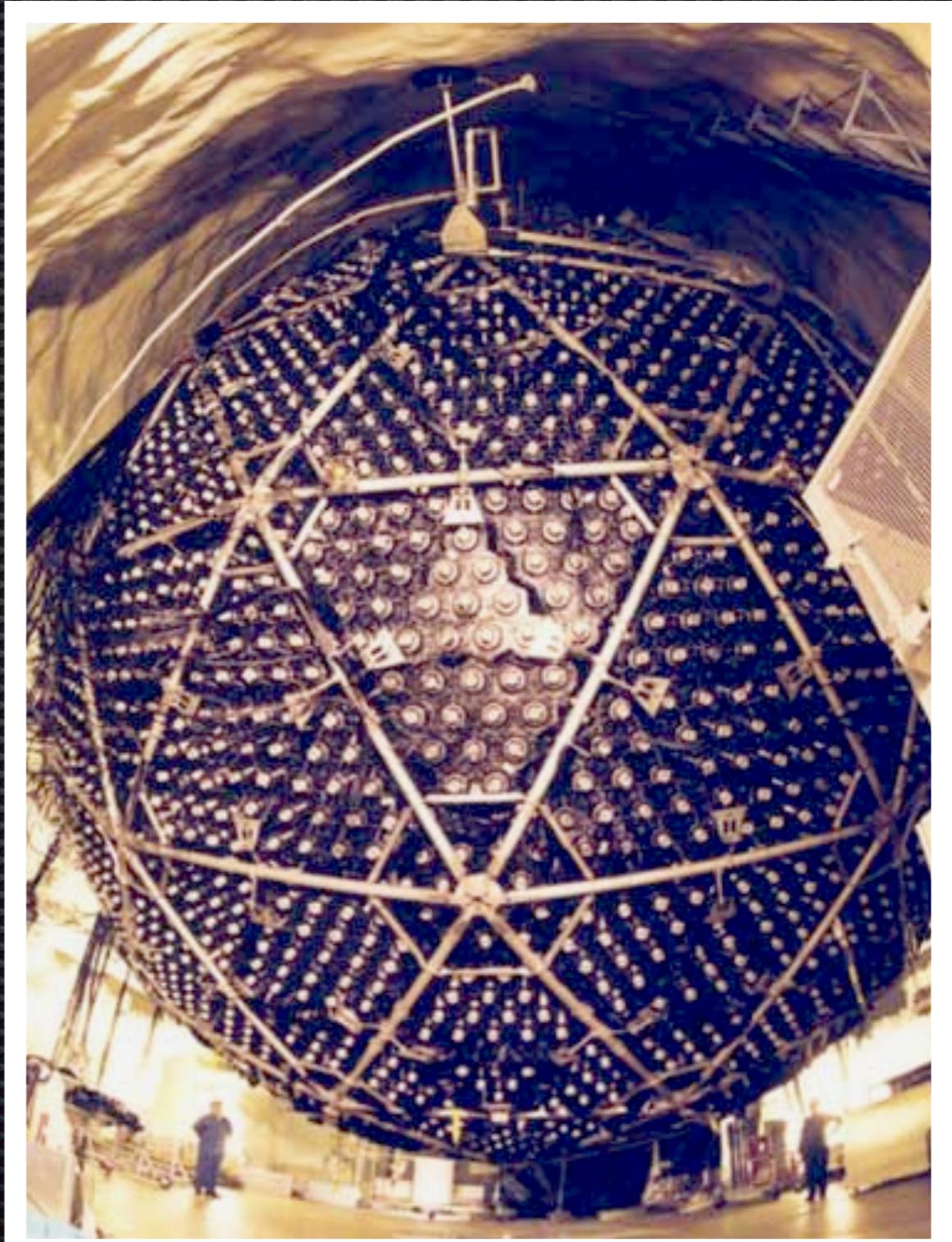
- 2000 Fermilab, DONUT experiment (Direct Observation of the Nu Tau)



Solar Oscillations Complete!

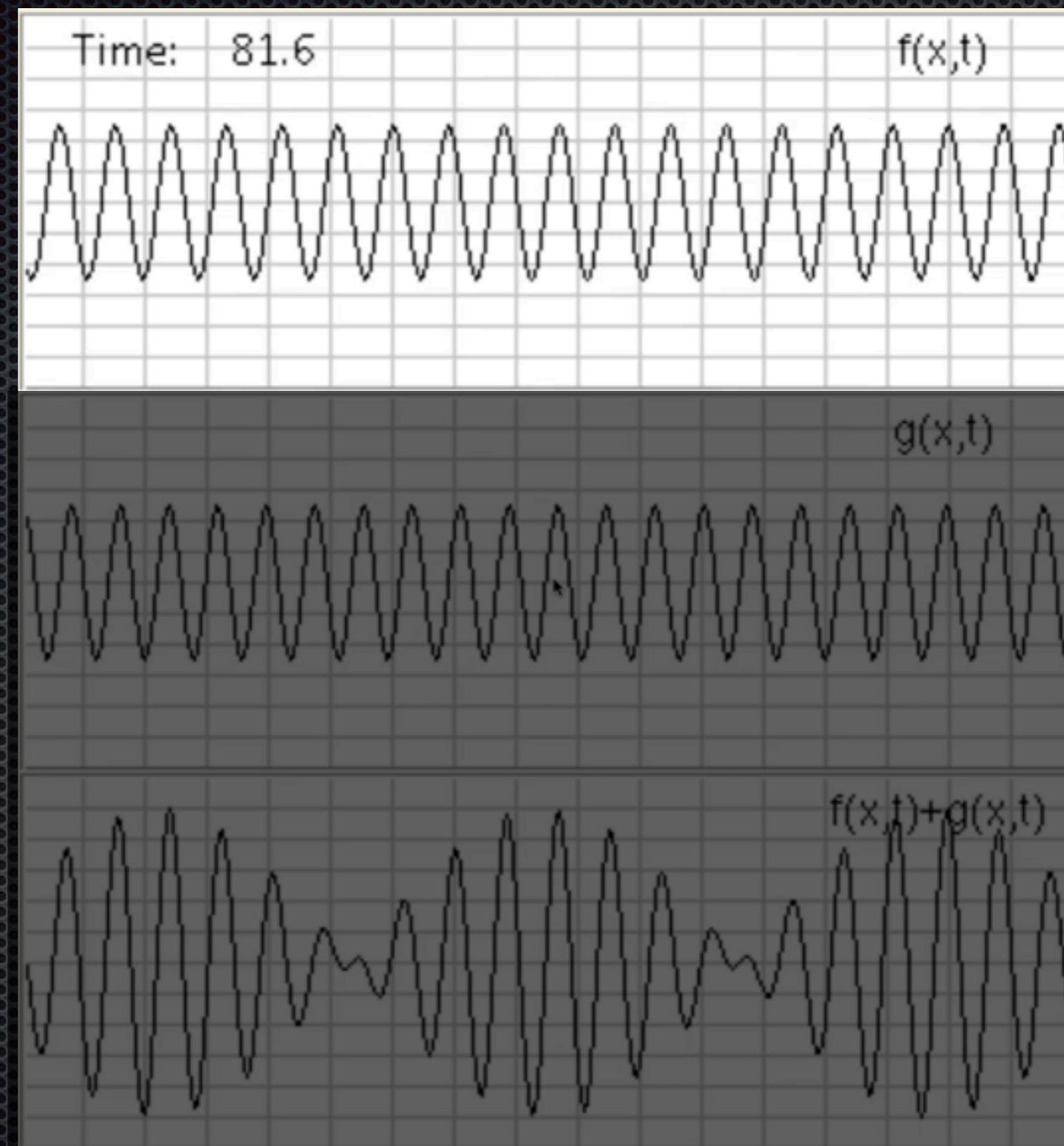
- 2001, SNO experiment (Sudbury Neutrino Observatory) in Canada.
- Heavy water cherenkov detector, sensitive to all three neutrino flavors.

ν_e sun $\rightarrow \nu_e, \nu_\mu, \nu_\tau$



Neutrino oscillation probability

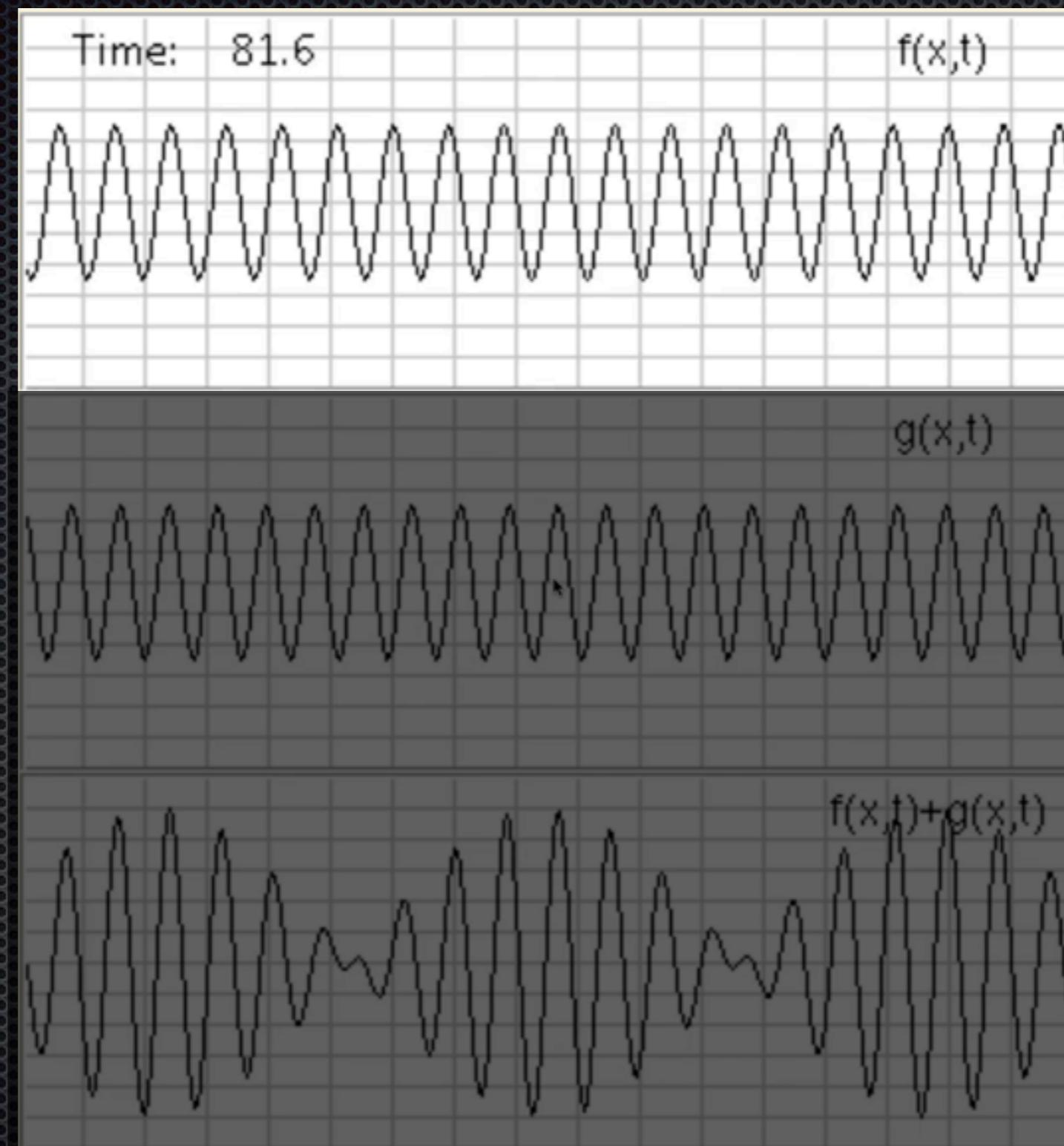
$\nu_e \rightarrow \nu_\mu$



440 Hz

Neutrino oscillation probability

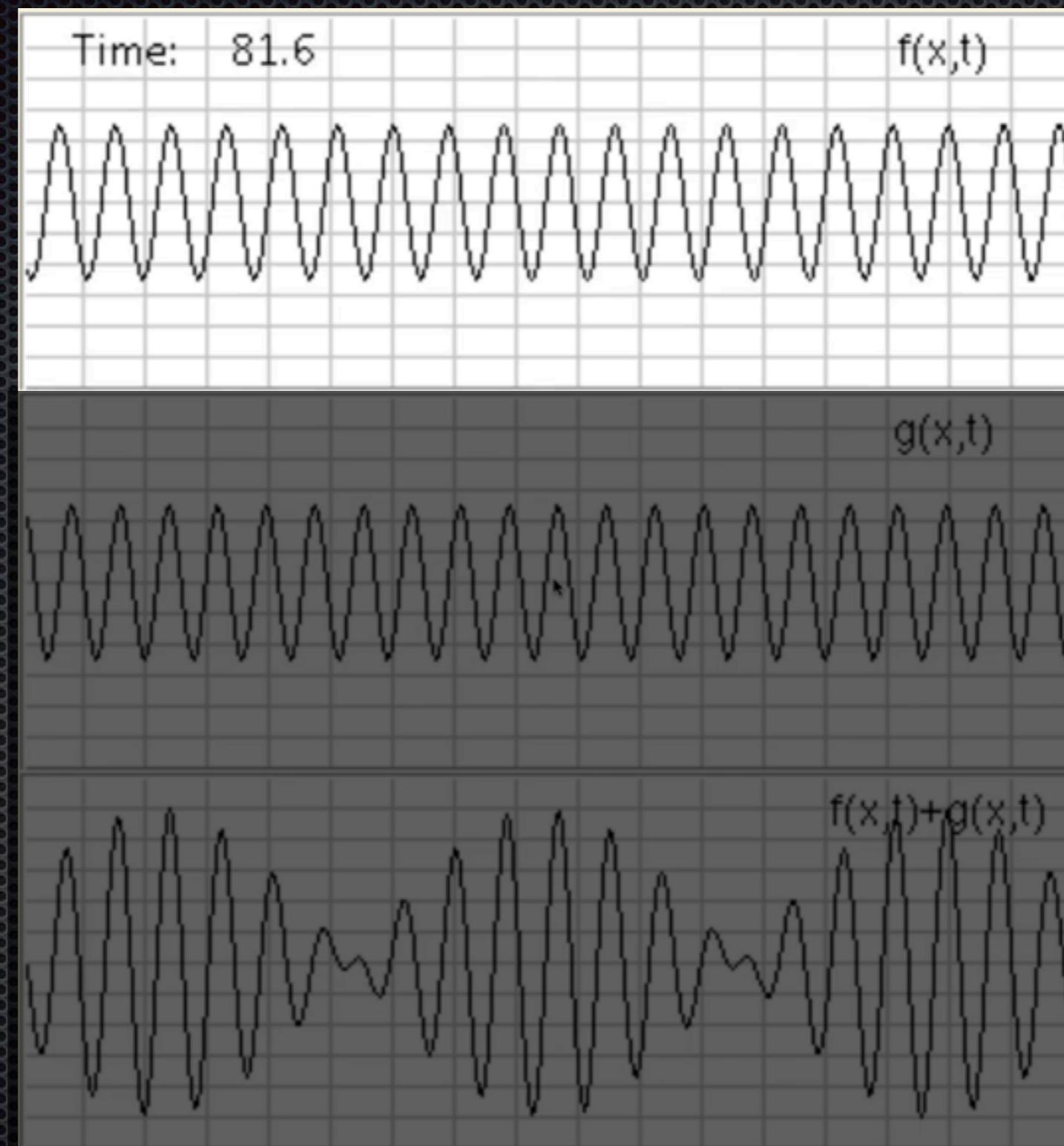
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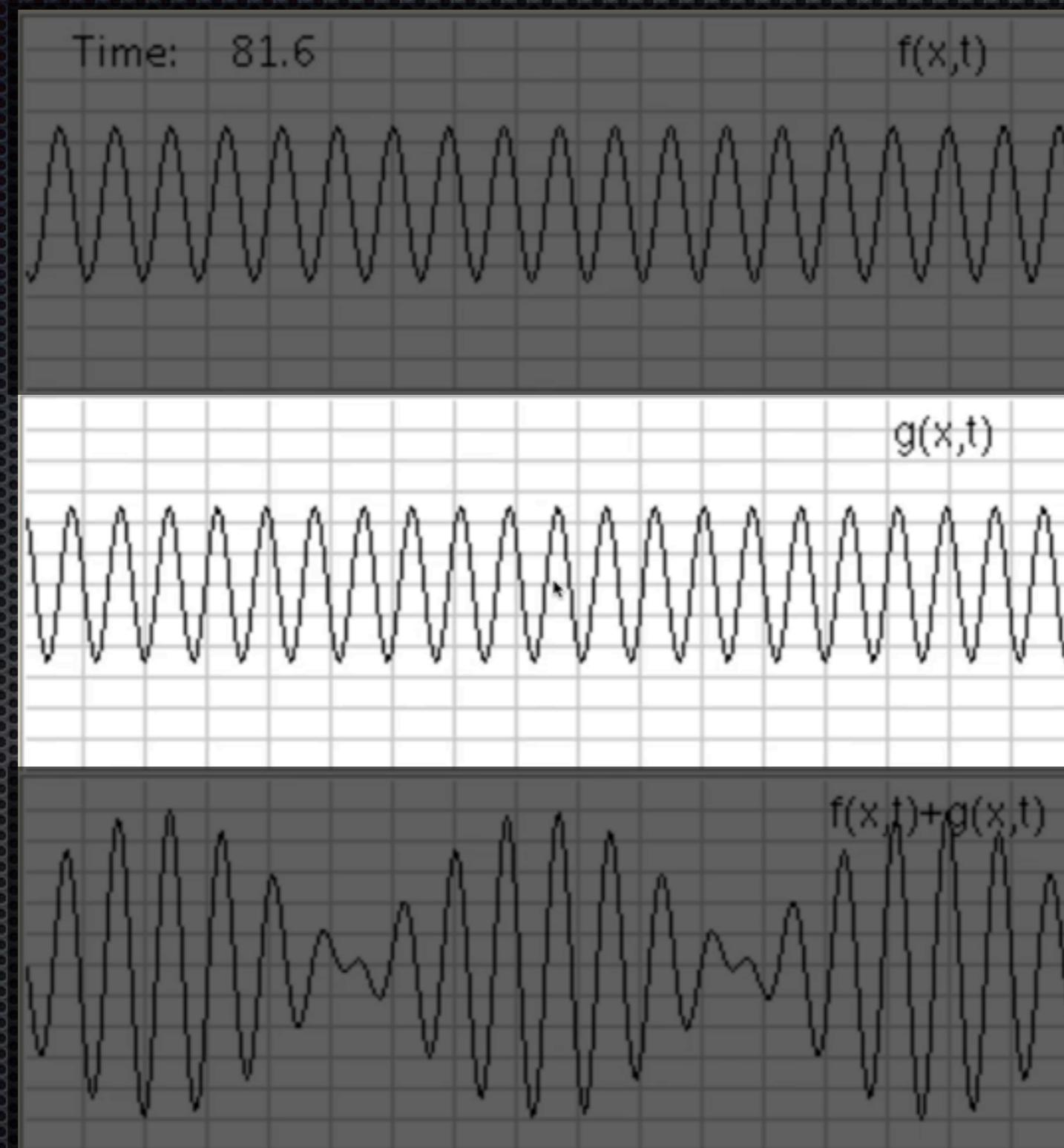
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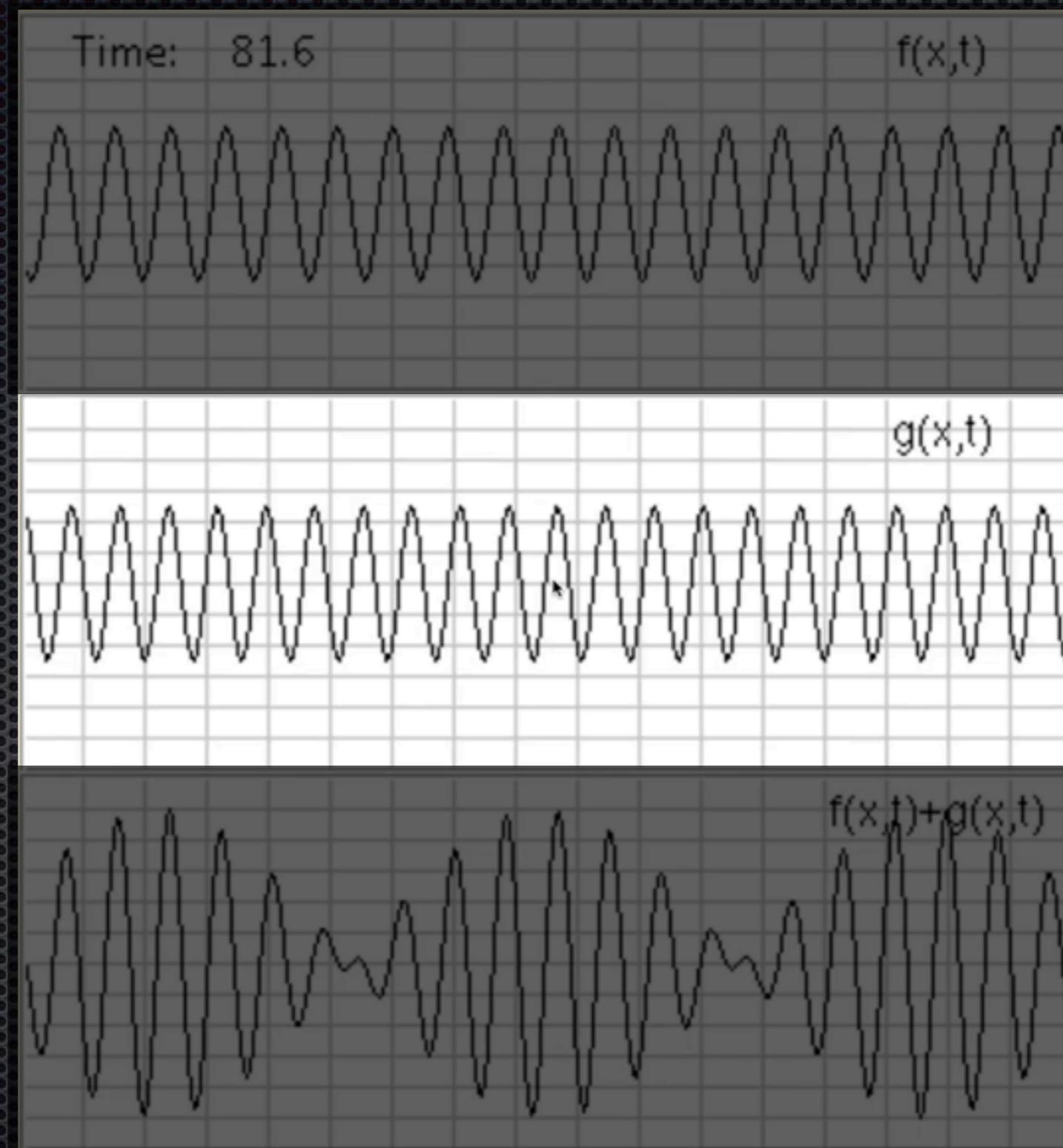
$\nu_e \rightarrow \nu_\tau$



441 Hz

Neutrino oscillation probability

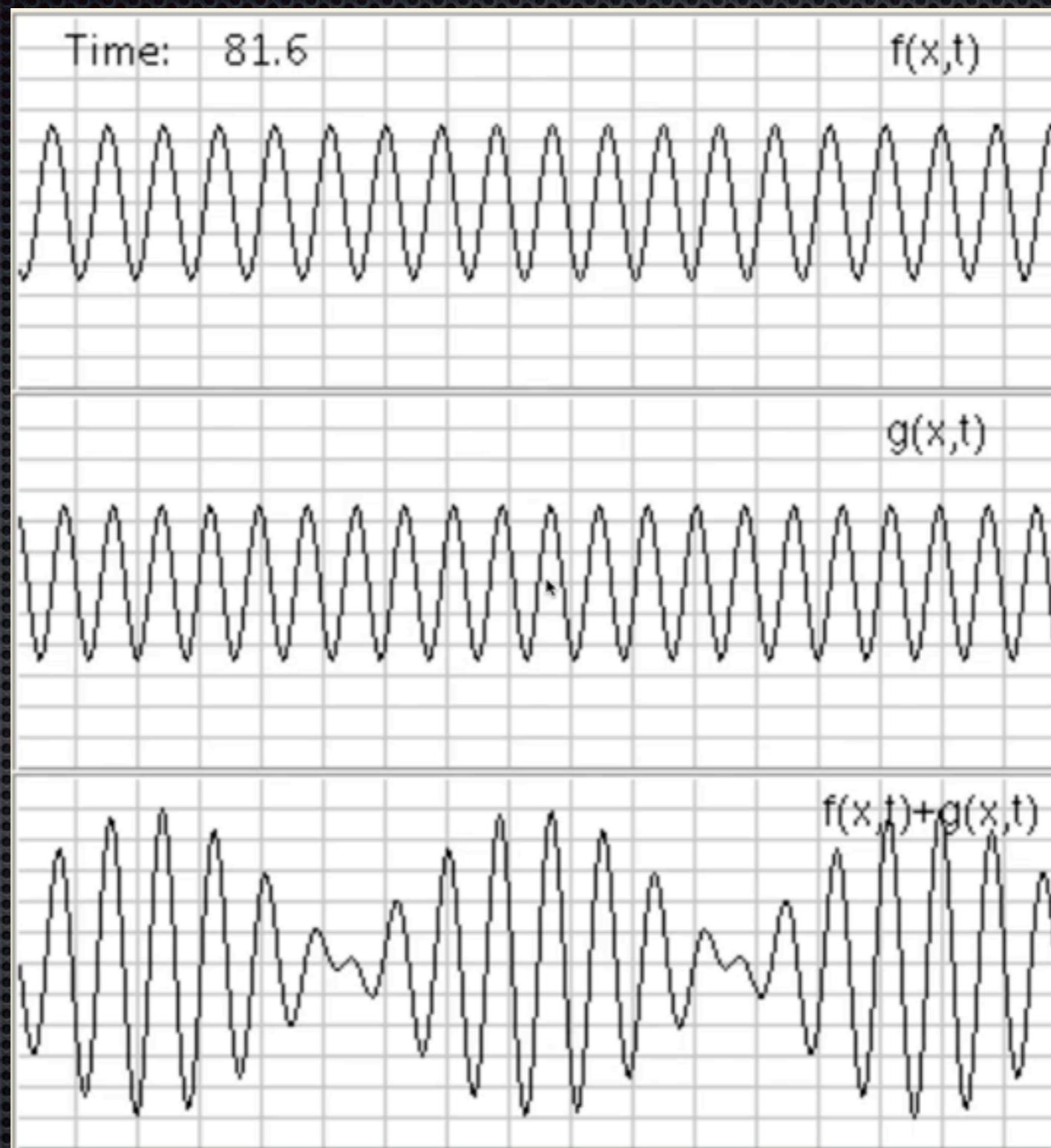
$\nu_e \rightarrow \nu_\tau$



441 Hz

Neutrino oscillation probability

$\nu_e \rightarrow \nu_\mu$



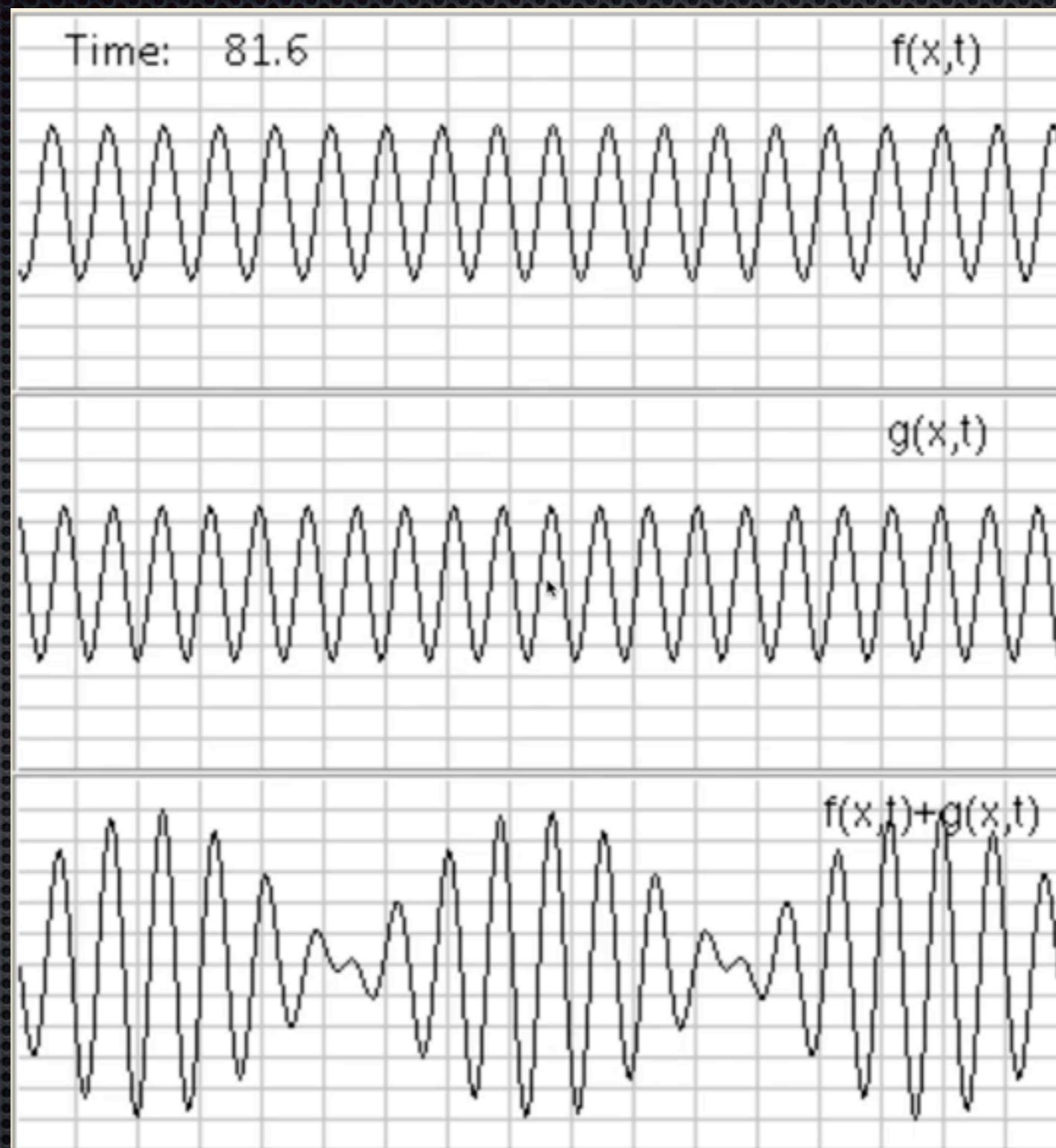
$\nu_e \rightarrow \nu_\tau$

$\nu_e \not\rightarrow \nu_e$

440 Hz + 441 Hz

Neutrino oscillation probability

$\nu_e \rightarrow \nu_\mu$



$\nu_e \rightarrow \nu_\tau$

$\nu_e \not\rightarrow \nu_e$

440 Hz + 441 Hz

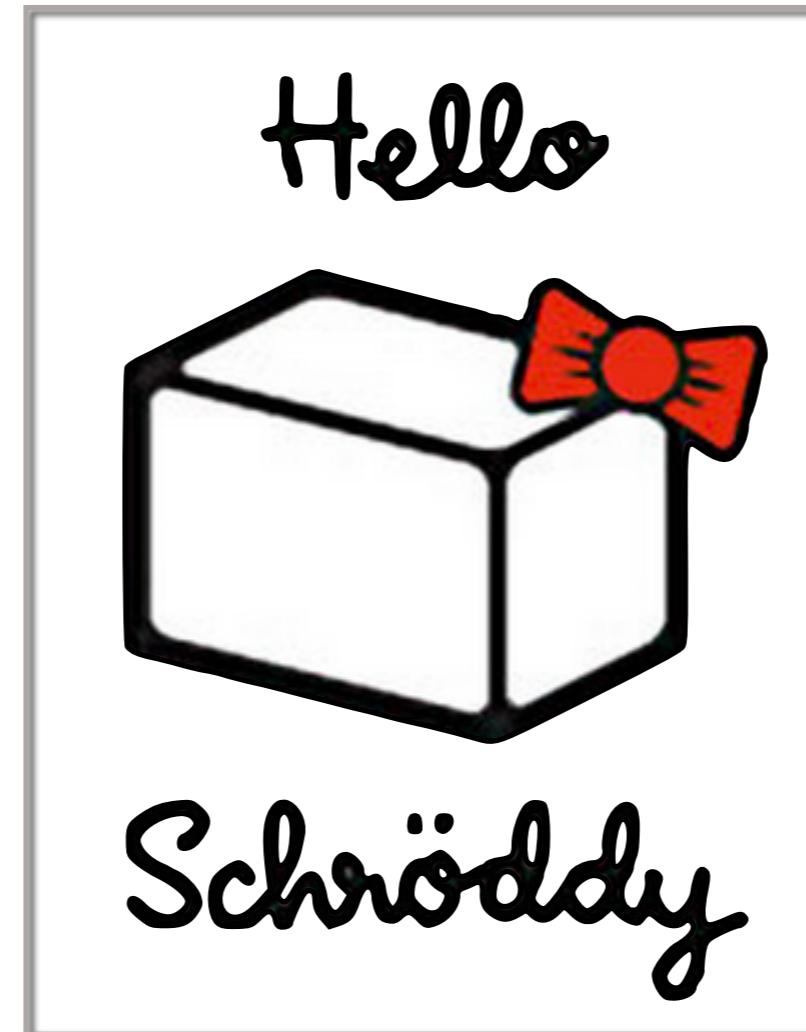
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Why do neutrinos oscillate?

- Review Schrödinger's Equation
- Free particle solution
- Assume some neutrino mixing
- See what happens later

Schrödinger's Equation

$$i\hbar \frac{\partial}{\partial t} \Psi(\mathbf{r}, t) = \hat{H}\Psi(\mathbf{r}, t)$$



Schrödinger's Equation

$$i\hbar \frac{\partial}{\partial t} \Psi(\mathbf{r}, t) = \hat{H}\Psi(\mathbf{r}, t)$$

Schrödinger's Equation

$$i\hbar \frac{\partial}{\partial t} \Psi(\mathbf{r}, t) = \hat{H} \Psi(\mathbf{r}, t)$$

$$\hat{H} = \hat{T} + \hat{V}$$

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$$\hat{H} = \hat{T} + \hat{V}$$



$$\hat{T} = \frac{\hat{\mathbf{p}} \cdot \hat{\mathbf{p}}}{2m} = \frac{\hat{p}^2}{2m} = -\frac{\hbar^2}{2m} \nabla^2$$

Schrödinger's Equation

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$$\hat{V} = V(\mathbf{r}, t)$$

Schrödinger's Equation

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$$i\hbar \frac{\partial}{\partial t} \Psi(\mathbf{r}, t) = -\frac{\hbar^2}{2m} \nabla^2 \Psi(\mathbf{r}, t) + V(\mathbf{r}, t) \Psi(\mathbf{r}, t)$$

Free Particle Solution

$$i\hbar \frac{\partial}{\partial t} \Psi(\mathbf{r}, t) = -\frac{\hbar^2}{2m} \nabla^2 \Psi(\mathbf{r}, t) + V(\mathbf{r}, t) \Psi(\mathbf{r}, t)$$

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- For free particle, let potential be zero
 - $V(r,t) = 0$

Free Particle Solution

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Free Particle Solution

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- $V(\mathbf{r}, t) = 0$

$$i\hbar \frac{\partial}{\partial t} \Psi(\mathbf{r}, t) = -\frac{\hbar^2}{2m} \nabla^2 \Psi(\mathbf{r}, t)$$

- Simplify laplacian, by selecting function to propagate in x direction: $\mathbf{r} = \hat{x}\mathbf{\hat{x}}$.

Free Particle Solution

$$i\hbar \frac{\partial}{\partial t} \Psi(\mathbf{r}, t) = -\frac{\hbar^2}{2m} \nabla^2 \Psi(\mathbf{r}, t) + V(\mathbf{r}, t) \Psi(\mathbf{r}, t)$$

- For free particle, let potential be zero

- $V(\mathbf{r}, t) = 0$

$$i\hbar \frac{\partial}{\partial t} \Psi(\mathbf{r}, t) = -\frac{\hbar^2}{2m} \nabla^2 \Psi(\mathbf{r}, t)$$

- Simplify laplacian, by selecting function to propagate in x direction: $\hat{\mathbf{r}} = \hat{x}$.

$$i\hbar \frac{\partial}{\partial t} \Psi(x, t) = -\frac{\hbar^2}{2m} \frac{\partial^2}{\partial x^2} \Psi(x, t)$$

Free Particle Solution

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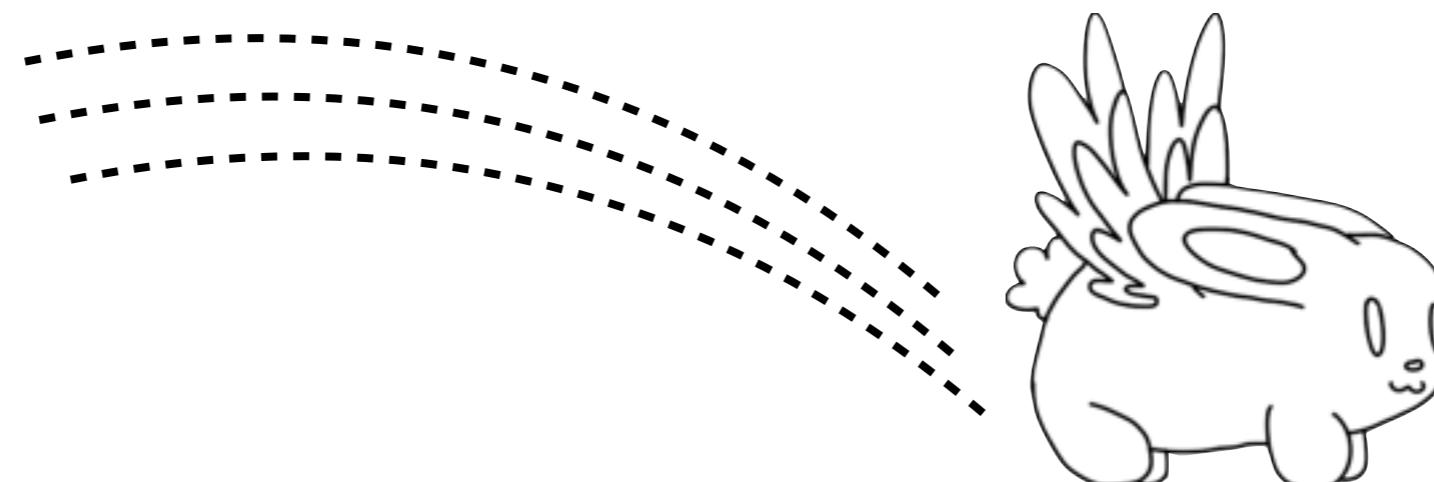
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- To simplify, let us consider the 2 neutrino mixing case.
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Assume flavor eigenstates can be written as a linear combination of mass eigenstates.

θ is the “mixing angle”

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Some tricks...

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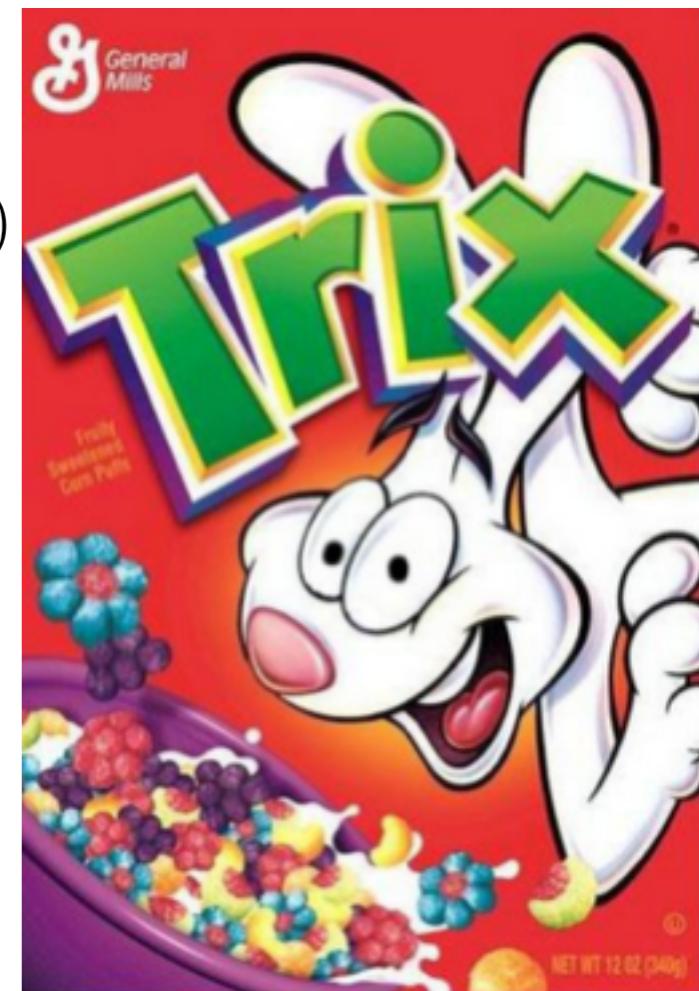
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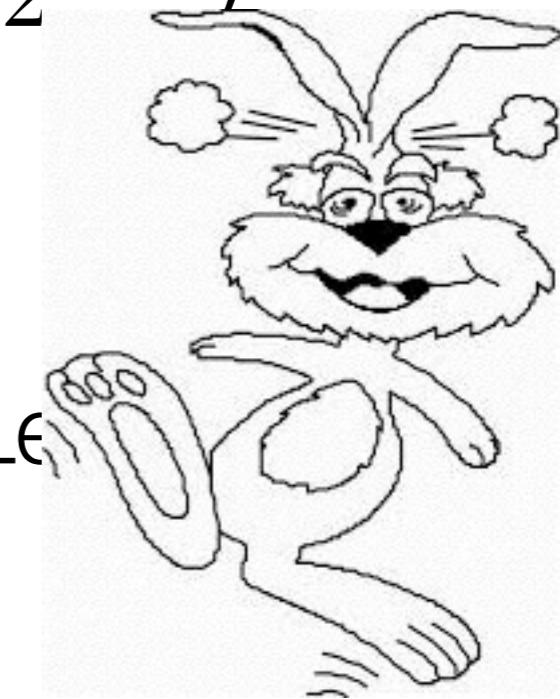
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Oscillation Probability

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where Δm_{12}^2 is in units of eV^2 , L is km, E is GeV

So what!?

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Wow! We just showed that oscillations imply that neutrinos have mass!

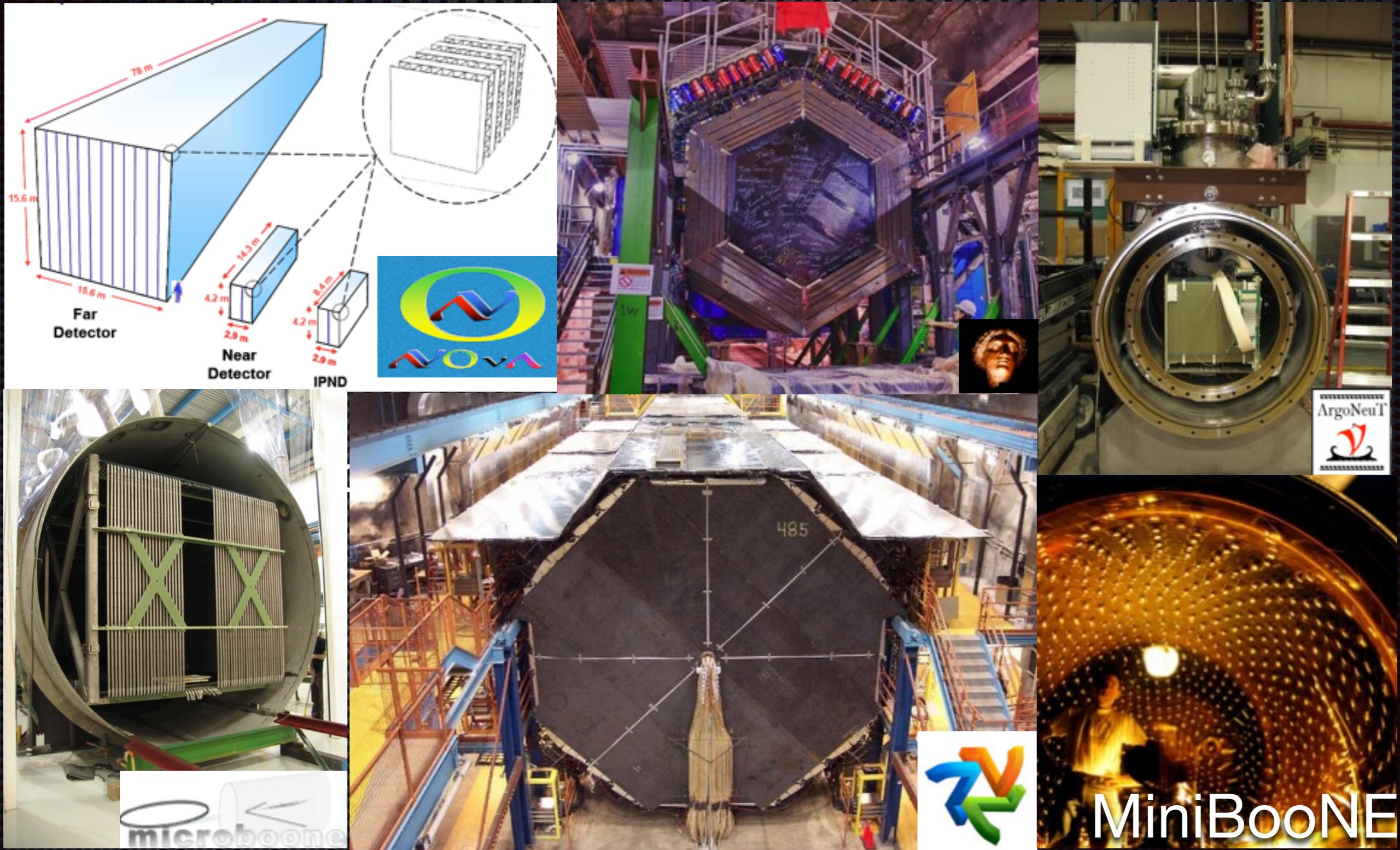
Still thirsty?

- It is left as an exercise to the reader to expand the 2 neutrino mixing case to the 3 neutrino case.



- ❖ What's particle physics? (translated into “Tia Speak”)
- ❖ Where are we in neutrino physics? (Story time!)
- ❖ FUN MATH for neutrino masses!
- ❖ Where are we going in neutrino physics? (A.K.A why Tia researches neutrinos)

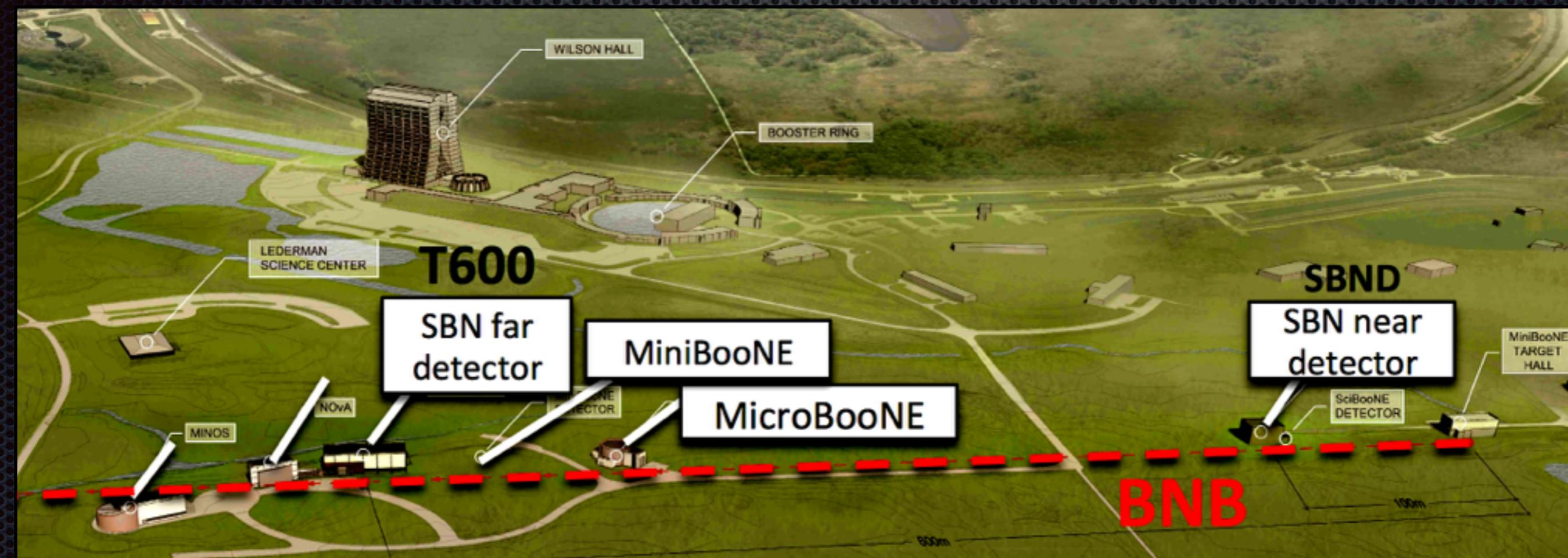
At Fermilab



At Fermilab

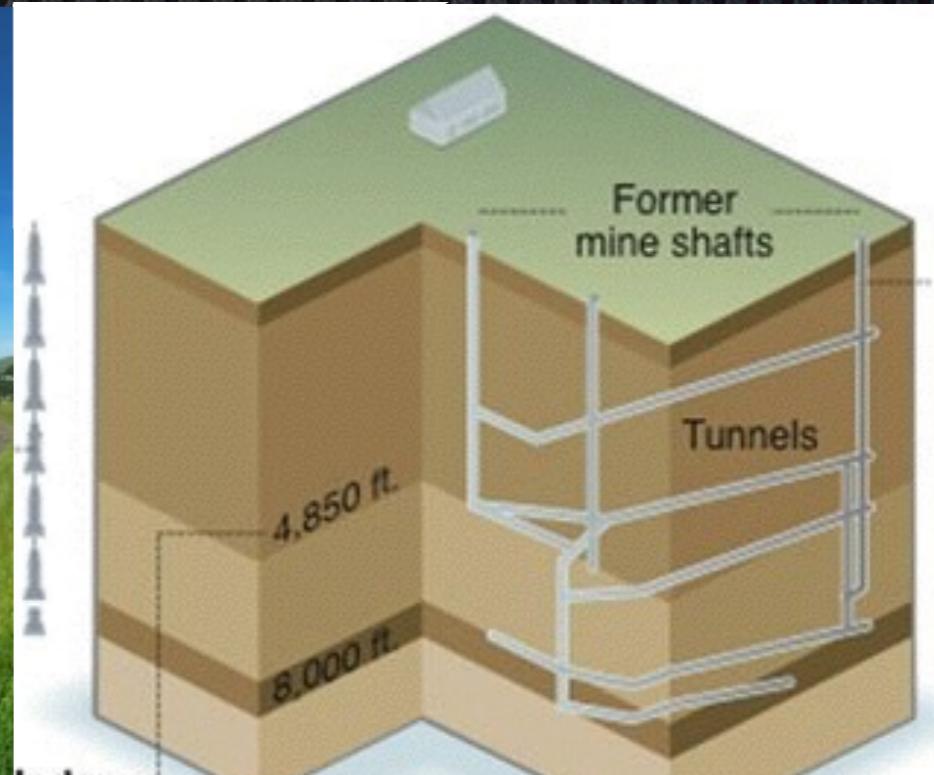


What the future holds for Fermilab



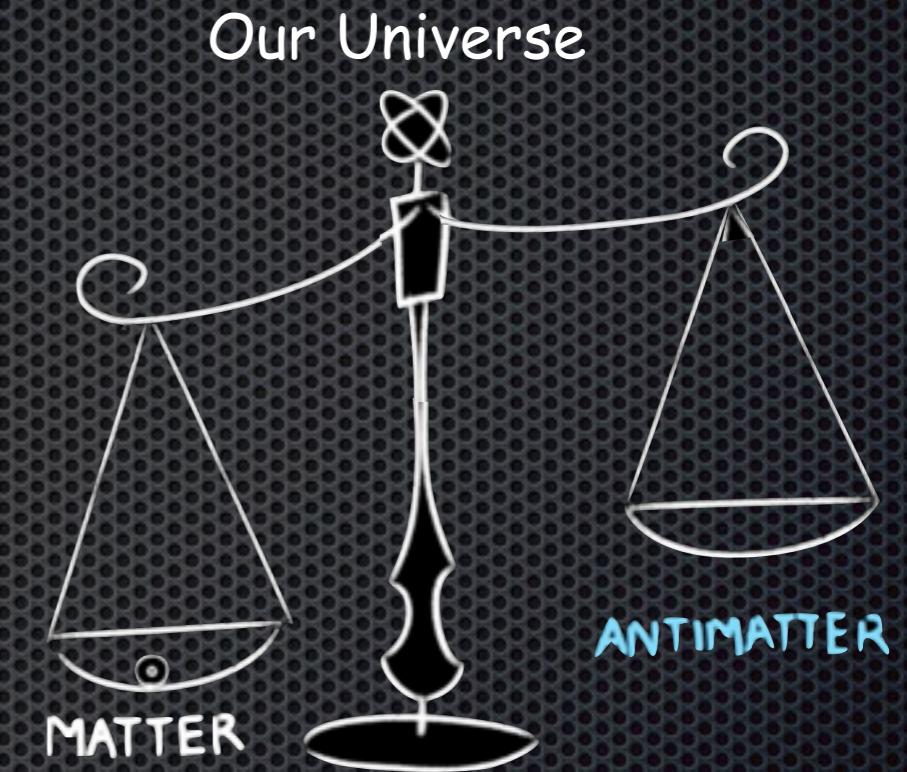
The Short Baseline Program

What the future holds for Fermilab



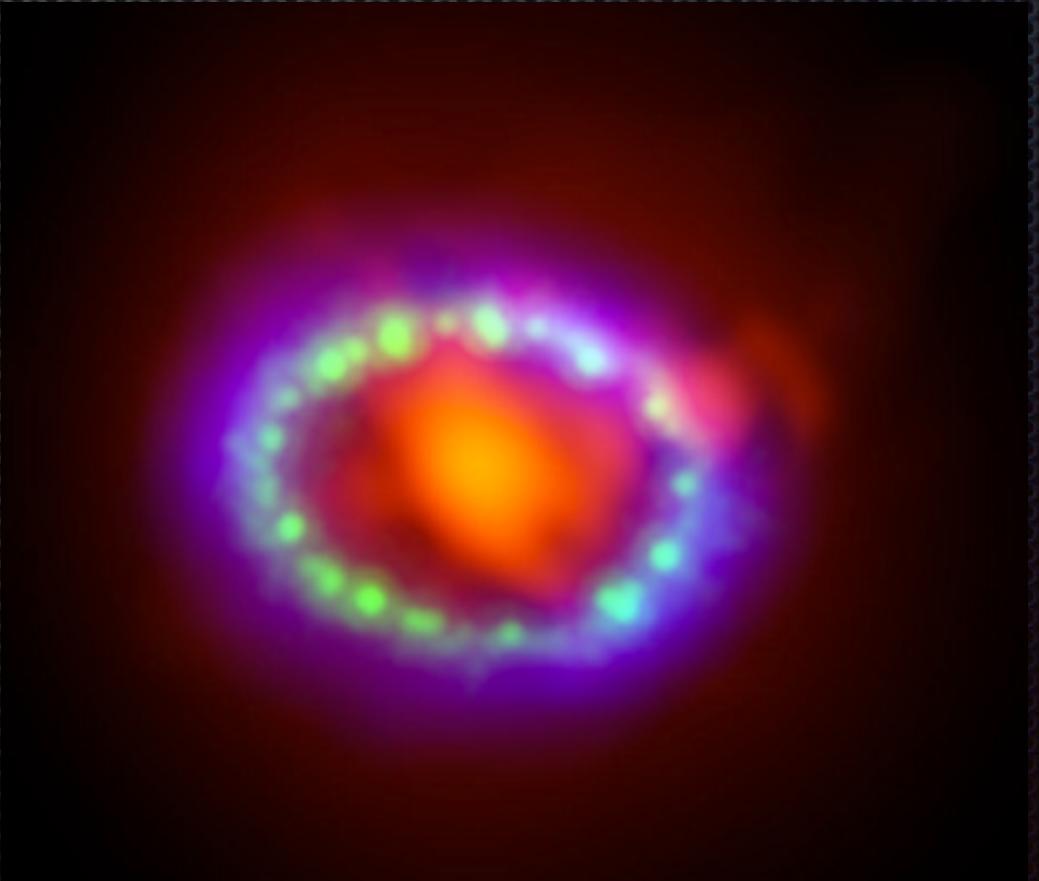
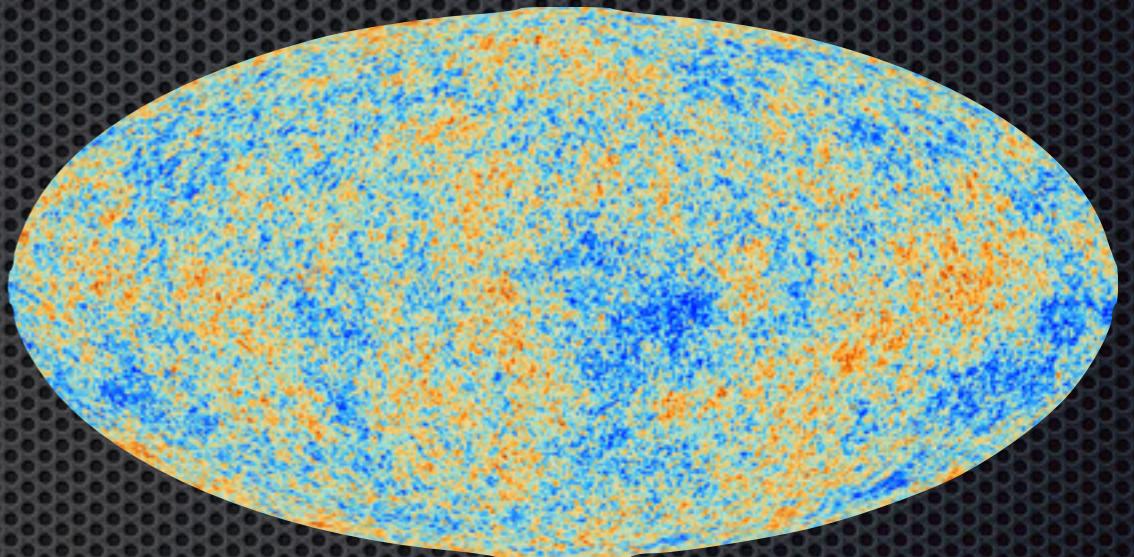
Why Tia researches neutrinos (1/3)

- Basic unknown properties!
 - Mass?
 - How many types?
 - Do neutrinos behave the same as anti-neutrinos?
 - neutrino = anti-neutrino?
“Majorana”



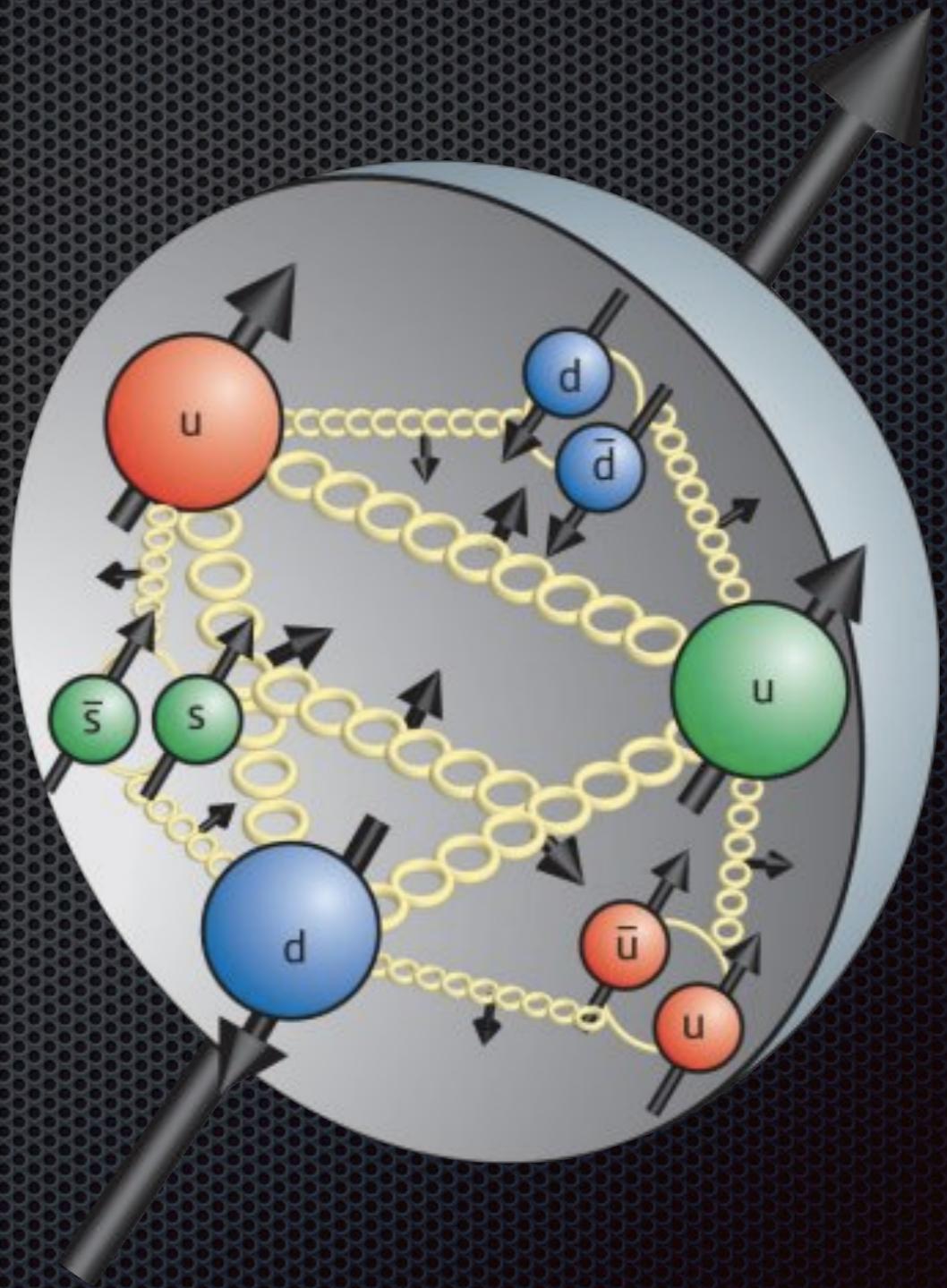
Why Tia researches neutrinos (2/3)

- Explore the universe
 - see where light can't! (Big bang relic neutrinos!?)
 - see inside stars! (solar, supernovae neutrinos)

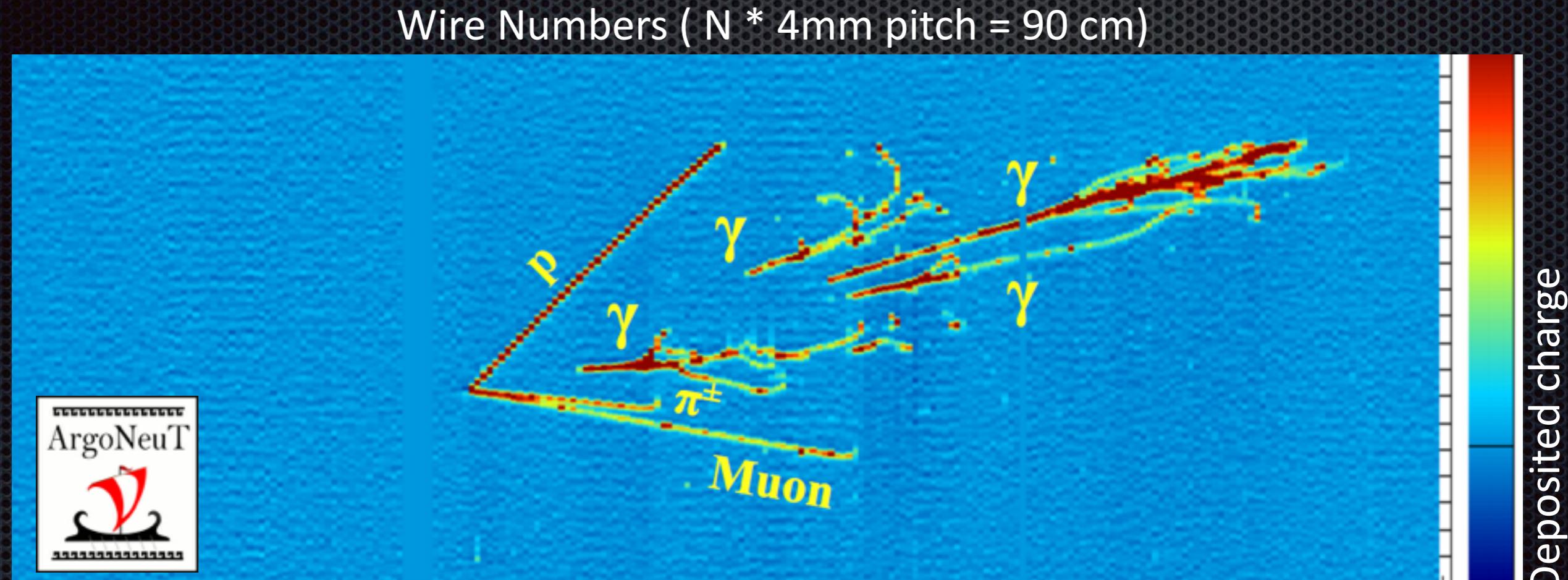


Why Tia researches neutrinos (3/3)

- Explore atoms
 - Neutrino scattering probabilities (cross-sections), used to refine experiments.
 - Spin composition of protons and neutrons. ❤

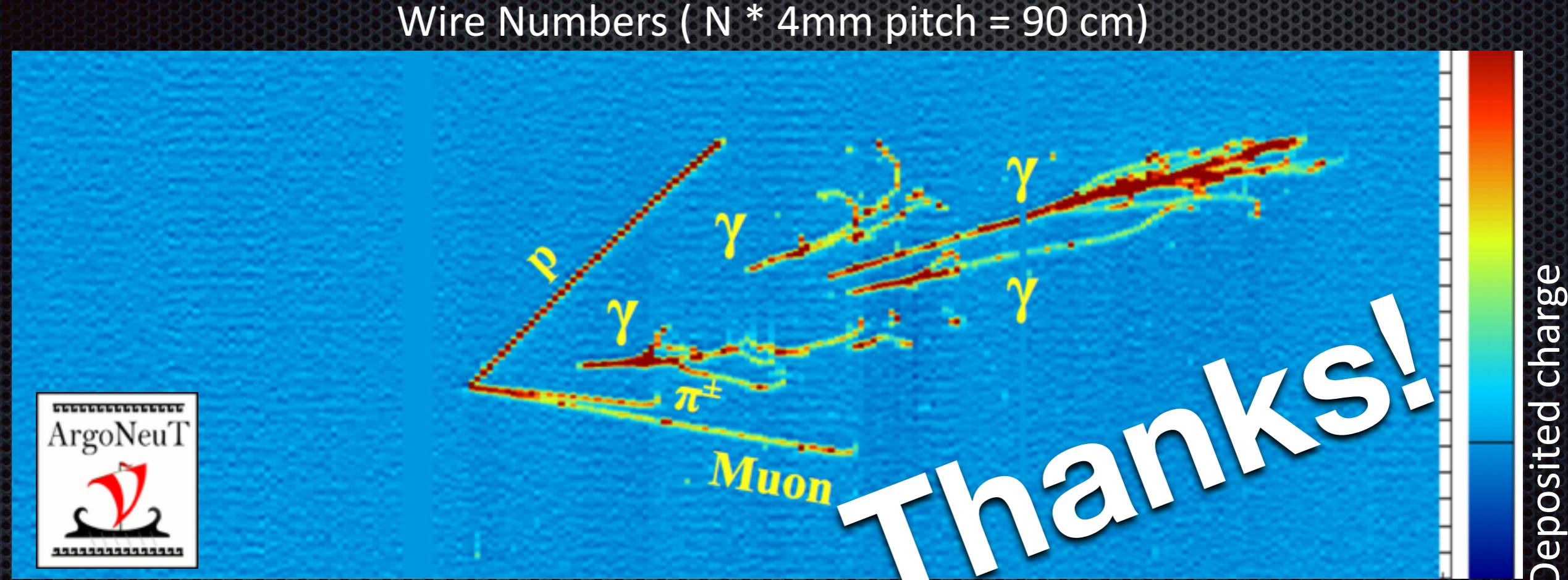


Amazing world of neutrinos!



What events will look like in MicroBooNE!
Commissioning now!
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